



X-shooter Spectroscopy and *HST* Imaging of 15 Massive Quiescent Galaxies at $z \gtrsim 2$

Mikkel Stockmann^{1,2,3} , Sune Toft^{1,2} , Anna Gallazzi⁴ , Stefano Zibetti⁴ , Christopher J. Conselice^{5,6} , Berta Margalef-Bentabol⁷ , Johannes Zabl⁸ , Inger Jørgensen⁹ , Georgios E. Magdis^{1,2,10} , Carlos Gómez-Guijarro^{1,2,3} , Francesco M. Valentino^{1,2,3} , Gabriel B. Brammer^{1,2} , Daniel Ceverino^{1,2} , Isabella Cortzen^{1,2,3} , Iary Davidzon^{1,2,11} , Richardo Demarco¹² , Andreas Faisst¹¹ , Michaela Hirschmann³ , Jens-Kristian Krogager¹³ , Claudia D. Lagos^{1,14,15} , Allison W. S. Man¹⁶ , Carl J. Mundy⁶ , Yingjie Peng¹⁷ , Jonatan Selsing^{1,2,3} , Charles L. Steinhardt^{1,2,3} , and Kathrine E. Whitaker^{18,19}

¹ Cosmic Dawn Center (DAWN), Denmark; mikkelstockmann@gmail.com

² Niels Bohr Institute, University of Copenhagen, Lyngbyvej 2, DK-2100 Copenhagen, Denmark

³ DARK, Niels Bohr Institute, University of Copenhagen, Lyngbyvej 2, DK-2100 Copenhagen, Denmark

⁴ INAF-Osservatorio Astrofisico di Arcetri, Largo Enrico Fermi 5, I-50125 Firenze, Italy

⁵ Centre for Astronomy and Particle Theory, School of Physics & Astronomy, University of Nottingham, Nottingham, NG7 2RD, UK

⁶ University of Nottingham, School of Physics & Astronomy, Nottingham, NG7 2RD, UK

⁷ LERMA, Observatoire de Paris, PSL Research University, CNRS, Sorbonne Universités, UPMC Univ. Paris 06, France

⁸ Univ Lyon, Univ Lyon1, Ens de Lyon, CNRS, Centre de Recherche Astrophysique de Lyon UMR5574, F-69230 Saint-Genis-Laval, France

⁹ Gemini Observatory, 670 N. A'ohoku Pl., Hilo, HI 96720, USA

¹⁰ Institute for Astronomy, Astrophysics, Space Applications and Remote Sensing, National Observatory of Athens, 15236, Athens, Greece

¹¹ IPAC, California Institute of Technology, 1200 E. California Blvd., Pasadena, CA 91125, USA

¹² Departamento de Astronomía, Facultad de Ciencias Físicas y Matemáticas, Universidad de Concepción, Concepción, Chile

¹³ Institut d'Astrophysique de Paris, UMR 7095, CNRS-SU, 98bis bd Arago, F-75014 Paris, France

¹⁴ International Centre for Radio Astronomy Research (ICRAR), M468, University of Western Australia, 35 Stirling Hwy., Crawley, WA 6009, Australia

¹⁵ ARC Centre of Excellence for All Sky Astrophysics in 3 Dimensions (ASTRO 3D), Australia

¹⁶ Dunlap Institute for Astronomy & Astrophysics, 50 St. George St., Toronto, ON M5S 3H4, Canada

¹⁷ Kavli Institute for Astronomy and Astrophysics, Peking University, 5 Yiheyuan Rd., Beijing 100871, People's Republic of China

¹⁸ Department of Physics, University of Connecticut, Storrs, CT 06269, USA

¹⁹ Department of Astronomy, University of Massachusetts, Amherst, MA 01003, USA

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Abstract

We present a detailed analysis of a large sample of spectroscopically confirmed massive quiescent galaxies (MQGs; $\log(M_*/M_\odot) \sim 11.5$) at $z \gtrsim 2$. This sample comprises 15 galaxies selected in the COSMOS and UDS fields by their bright *K*-band magnitudes and followed up with Very Large Telescope (VLT) X-shooter spectroscopy and *Hubble Space Telescope* (*HST*)/WFC3 H_{F160W} imaging. These observations allow us to unambiguously confirm their redshifts, ascertain their quiescent nature and stellar ages, and reliably assess their internal kinematics and effective radii. We find that these galaxies are compact, consistent with the high-mass end of the stellar mass–size relation for quiescent galaxies at $z = 2$. Moreover, the distribution of the measured stellar velocity dispersions of the sample is consistent with the most massive local early-type galaxies from the MASSIVE Survey, showing that evolution in these galaxies is dominated by changes in size. The *HST* images reveal, as surprisingly high, that 40% of the sample has tidal features suggestive of mergers and companions in close proximity, including three galaxies experiencing ongoing major mergers. The absence of velocity dispersion evolution from $z = 2$ to 0, coupled with a doubling of the stellar mass, with a factor of 4 size increase and the observed disturbed stellar morphologies, supports dry minor mergers as the primary drivers of the evolution of the MQGs over the last 10 billion yr.

Unified Astronomy Thesaurus concepts: Infrared galaxies (790); Compact galaxies (285); Quenched galaxies (2016); Near infrared astronomy (1093); Infrared astronomy (786); Galaxy kinematics (602); Galaxy dynamics (591); Stellar masses (1614); Stellar ages (1581); Stellar radii (1626); High resolution spectroscopy (2096); Hubble Space Telescope (761)

1. Introduction

Local galaxies follow a bimodal distribution in color represented by blue star-forming spirals and red dormant elliptical galaxies. The most massive galaxies, primarily located in cluster environments, are the giant elliptical galaxies with stellar population ages suggesting formation more than 10 billion yr ago (Ma et al. 2014; Greene et al. 2015).

A population of red massive galaxies was discovered to exist at $z \sim 2$ (Franx et al. 2003; Daddi et al. 2004) and subsequently confirmed to have a quiescent stellar population (Cimatti et al. 2004; Daddi et al. 2005; Labbé et al. 2005; Kriek et al. 2006a; Toft et al. 2007; Williams et al. 2009). At this epoch, the star

formation rate (SFR) density peaked (Madau & Dickinson 2014) alongside substantial active galactic nuclei (AGNs; Hopkins et al. 2007). At this time, half of the most massive ($\log_{10}(M_*/M_\odot) > 11$) galaxies were already devoid of star formation (SF) and had old stellar ages, suggesting that they quenched their SF at even earlier times ($z > 3$), when the universe was only a few Gyr old (e.g., Kriek et al. 2006b; van Dokkum et al. 2006, 2008; Franx et al. 2008; Toft et al. 2009; McCracken et al. 2010; Williams et al. 2010; Brammer et al. 2011; Whitaker et al. 2011; Wuyts et al. 2011; Kado-Fong et al. 2017; Morishita et al. 2018). Nowadays, quiescent galaxies are popularly defined by the *UVJ* color–color relations (see, e.g., Muzzin et al. 2013a).

These massive quiescent galaxies (MQGs) are found to be remarkably compact with extremely high stellar densities when compared to local galaxies with similar stellar mass (Papovich et al. 2005; Trujillo et al. 2006, 2007; Buitrago et al. 2008; Cimatti et al. 2008; van Dokkum et al. 2008; Bezanson et al. 2009; Conselice et al. 2011; Szomoru et al. 2012; van der Wel et al. 2014; Mowla et al. 2018). A small number of elliptical galaxies this compact are found in the local universe (Trujillo et al. 2009; Taylor et al. 2010b; Shih & Stockton 2011; Ferré-Mateu et al. 2012), but these are too young (ages $\sim 2\text{--}4$ Gyr) to be the descendants of $z = 2$ compact quiescent galaxies. This suggests that the vast majority of the $z = 2$ population must undergo a substantial increase in size to evolve into local elliptical galaxies (Bell et al. 2012).

Bluck et al. (2012) found that the expected size evolution between $z = 2.5$ and the present day can be described primarily by minor mergers. However, Newman et al. (2012) and Man et al. (2016b) found that minor mergers can account for the evolution at $z < 1$ and that additional mechanisms of growth are required at higher redshift. The minor merger scenario is supported by the continuous size evolution found in the compilation of spectroscopic (Damjanov et al. 2011; Belli et al. 2014b; Matharu et al. 2019) and photometric (van der Wel et al. 2014; Faisst et al. 2017; Mowla et al. 2018) studies, as well as the expected theoretical predictions of the galaxy properties during merger evolution (e.g., Khochfar & Silk 2006; Naab et al. 2009; Lagos et al. 2018).

To study the dynamics of MQGs at $z > 2$, it is important to obtain both reliable kinematic and morphological measurements using deep spectroscopic observations and high-resolution (adaptive optics or space-based) imaging (Kriek et al. 2009; Toft et al. 2012; van de Sande et al. 2013; Belli et al. 2017). Quiescent galaxies beyond $z > 2$ are more disklike with higher ellipticities than local ellipticals (Toft et al. 2005, 2007; van der Wel et al. 2011; Wuyts et al. 2011), which may cause heightened dispersion measurements from the contribution of unresolved rotation. In Toft et al. (2017) and Newman et al. (2018), the first spatially resolved gravitationally lensed $z > 2$ MQGs are observed.

At $z \sim 2$, MQGs are rare (Arcila-Osejo et al. 2019), and their quiescent nature implies faint rest-frame UV continua with no strong emission lines. Due to their rarity, large survey fields are essential to locate these galaxies. So far, only a small sample of MQGs have been spectroscopically confirmed at $z > 2$, in existing surveys like CANDELS+GOODS, and few of those have robust velocity dispersion measurements (van de Sande et al. 2013; Belli et al. 2014b, 2017; Kriek et al. 2016; Morishita et al. 2018).

In this paper, the structural and dynamical properties of 15 UVJ MQGs, $\log_{10}(M_*/M_\odot) > 11$, at $z > 2$ are studied, doubling the spectroscopically confirmed and absorption line-detected sample at this epoch using the 2 deg² Cosmic Evolution Survey (COSMOS) and Ultra Deep Survey (UDS) fields. These MQGs are examined in detail through their evolution to local galaxies and how they likely formed in minor and major merger processes. In a follow up paper (M. Stockmann et al. 2020, in preparation), the fundamental plane relation and its evolution to $z = 0$ is studied (Djorgovski & Davis 1987; Dressler et al. 1987).

In Section 2, we present the sample selection of the $z = 2$ galaxies and a corresponding local reference sample. The X-shooter spectroscopic and *Hubble Space Telescope* (HST)

imaging data reduction, alongside the photometry used throughout the paper, are presented in Section 3. In Section 4 we present the methods used to extract the X-shooter absorption line kinematics and stellar populations and HST structural properties from the data, together with a multi-wavelength comparison of different SF tracers. We address the issue of progenitor bias using our local reference sample in Section 5.1. We present the stellar population and kinematic and structural results in Sections 5.2 and 5.3 and the dynamical properties in Section 5.4. The results and evolution of these galaxies to $z = 0$ are discussed and summarized in Sections 6 and 7, respectively.

Throughout the manuscript, magnitudes are quoted in the AB system (Oke & Gunn 1983; Fukugita et al. 1996), and the following cosmological parameters are used: $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. All stellar masses are presented using the Chabrier (2003) initial mass function (IMF).

2. Sample Selection

The sample studied here consists of 15 MQGs from the COSMOS and UDS (Williams et al. 2009) fields for spectroscopic follow-up and is selected based on the modeling of their optical-to-far-infrared (FIR) broadband spectral energy distributions (SEDs). Three samples from three periods of observation are presented below. In the first program, galaxies were identified to be at $z_{\text{phot}} > 1.6$ and with old (> 1 Gyr) quiescent stellar populations (specific SFRs $\log(\text{sSFR}/\text{yr}) < -11$) in the updated version of the Ilbert et al. (2009) catalog of the COSMOS field described in Man et al. (2012). The four *K*-band-brightest ($K < 21.5$) sources covered by parallel HST/NICMOS observations were selected for follow-up to enable study of their morphology. These galaxies are referred to as the P86 sample, named after the period of Very Large Telescope (VLT)/X-shooter observations (P86; 2010–2011).

In a second program, 10 of the *K*-band-brightest ($K < 20.5$) galaxies in the COSMOS field with photometric redshifts²⁰ $z_{\text{phot}} > 1.9$, sSFRs $\log(\text{sSFR}/\text{yr}) < -10$, and stellar masses $\log_{10}(M_*/M_\odot) > 11$ from the Muzzin et al. (2013a) catalog were selected for follow-up. Based on visual inspection, the sources with nearby bright objects in the *K*-band images are excluded to avoid photometric contamination. Objects with *Spitzer*/MIPS 24 μm detections are also excluded to avoid either dusty star-forming galaxies or AGNs (Le Floch et al. 2009). Their SEDs were visually inspected, and galaxies with noisy photometry or bad fits were excluded. This pool of galaxies is dubbed the P93 sample, observed 3 yr after P86.

Finally, in the analysis presented here, the MQG UDS 19627 from Toft et al. (2012) is included. This object is selected as part of early VLT/X-shooter GTO observations to be quiescent ($\log(\text{sSFR}/\text{yr}) < -10$) at a high redshift ($z_{\text{phot}} = 2.02^{+0.07}_{-0.08}$) and a bright source ($K = 20.19$) in the UKIRT UDS (Williams et al. 2009). New HST/WFC3 H_{F160W} imaging of this galaxy is presented, allowing us to measure resolved morphology. The MQG UDS 19627 is minimally gravitationally lensed, but Toft et al. (2012) showed that, after taking this effect into account, the systematic change in magnification factor of 10%–20% correspond to 0.07 and 0.03 dex, resulting in lower stellar and dynamical mass.

²⁰ Using redshift quality parameters with odds = 1 from Brammer et al. (2008).

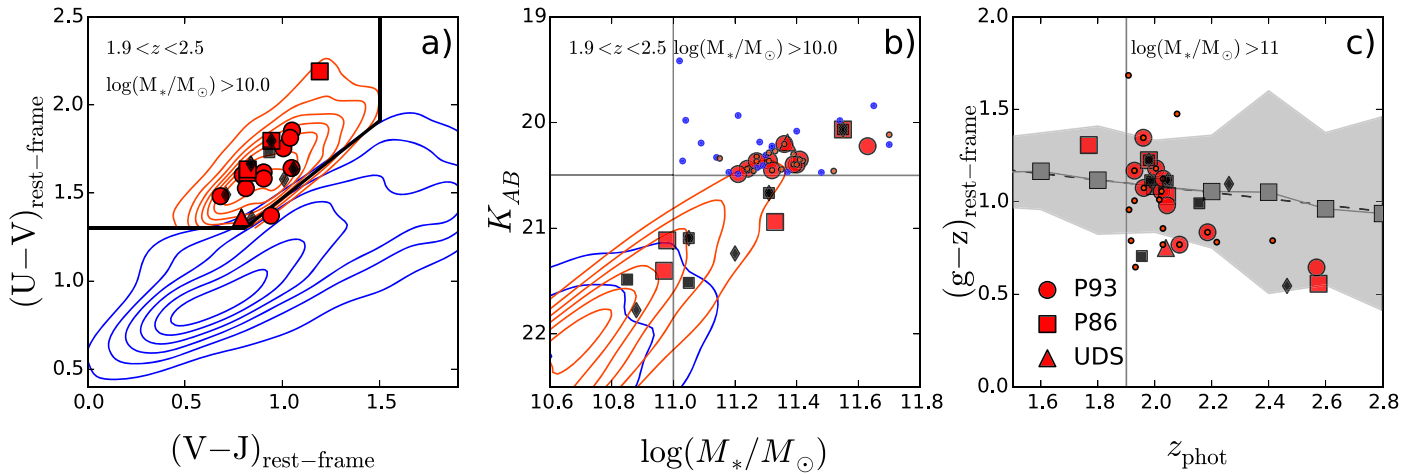


Figure 1. Photometric properties of the galaxy sample (red symbols; see legend in right panel) in the UVJ (a), K_{AB} - $\log(M_*/M_\odot)$ (b), and z_{phot} -rest-frame ($g-z$) (c) planes from the Muzzin et al. (2013a) catalog. Note that for UDS 19627, we use the Toft et al. (2012) K -band, stellar mass, z_{phot} , and rest-frame colors estimated from the observed photometry with EAZY (Brammer et al. 2011). The UVJ quiescent (red) and star-forming (blue) galaxies are shown by contours in the range $1.9 < z_{\text{phot}} < 2.5$ and $\log(M_*/M_\odot) > 10$ (Muzzin et al. 2013a). The spectroscopically confirmed $z > 2$ MQGs from COSMOS are shown with black symbols (squares, Krogager et al. 2014; diamonds, Belli et al. 2017). The small red/blue points in panel (b) are the galaxies that satisfy the criteria $K < 20.5$ and $\log(M_*/M_\odot) > 11$. The gray squares in panel (c) represent the running mean of the rest-frame ($g-z$) color of the massive, $\log_{10}(M_*/M_\odot) > 11$, UVJ -selected quiescent galaxies with the 1σ standard deviation in gray.

Our full sample is compiled from the three presented subgroups selected with variations in criteria on stellar mass, sSFR, and K -band brightness. In Figure 1(a), we show that despite the variation in selection criteria, this sample populates the quiescent galaxy region of the UVJ rest-frame color-color diagram (Muzzin et al. 2013b). For the sake of homogeneity, the full sample (except for UDS 19627) is shown using the Muzzin et al. (2013a) catalog. Our galaxies are consistent with the UVJ selection for massive ($\log(M_*/M_\odot) > 10$) quenched objects at $1.9 < z < 2.5$.

Figure 1(b) shows the position of our sample in the K -band magnitude-stellar mass plane. The $K < 20.5$ and $\log(M_*/M_\odot) > 11$ selection of the P93 sample results in significantly larger stellar masses than the average for the P86 sample (selected as MQGs with NICMOS coverage), with only one galaxy from the latter fully satisfying the criteria of P93 (previously presented in, among others, van de Sande et al. 2013; Kriek et al. 2016; Belli et al. 2018). The power of adding a minimum K -band threshold to the stellar mass criterion to select the most extreme MQGs is evident when comparing our sample with previous studies (van de Sande et al. 2013; Krogager et al. 2014; Belli et al. 2017), identifying, on average, MQGs with lower stellar masses. Our sample represents 60% of the total number of UVJ MQGs (29% of all galaxies) at $1.9 < z < 2.5$, $\log(M_*/M_\odot) > 11$, and $K < 20.5$ from Muzzin et al. (2013a; upper right corner of Figure 1(b)). We confirm that our selection of UVJ quiescent galaxies can be considered representative of the massive and K -band-brightest galaxies at $1.9 < z < 2.5$. This is done by using a modified version of the Anderson-Darling test²¹ to compare our stellar mass and K -band selection with the photometric samples.

One concern addressed by van de Sande et al. (2014) is that the selection of the K -band-brightest galaxies introduces a bias toward the bluest galaxies in the rest-frame color ($g-z$)_{rf}. To address this issue, the rest-frame colors ($g-z$)_{rf}, as a function of redshift between our sample and the UVJ -selected massive

($\log(M_*/M_\odot) > 11$) quiescent galaxies from Muzzin et al. (2013a), are compared in Figure 1(c). Contrary to the sample of van de Sande et al. (2014), 13/15 of our galaxies have ($g-z$)_{rf} colors consistent within the standard deviation of the average MQGs at a matching epoch. The Anderson-Darling test for k -samples confirms that the ($g-z$)_{rf} colors for our MQGs are representative of the ($\log(M_*/M_\odot) > 11$) UVJ MQGs at $1.9 < z < 2.5$. This suggests that our K -band-selected sample is, on average, not biased toward galaxies with bluer colors. However, the highest-redshift sources have systematic lower ($g-z$)_{rf} colors and could be subjected to this selection bias.

In summary, our sample is selected to be the most massive K -band-bright UVJ quiescent galaxies at $z > 2$. The selection is not subjected to a bias in ($g-z$)_{rf} and can be considered a 60% stellar mass and K -band complete sample of the quiescent galaxies at $z > 2$.

2.1. A Suitable Reference Sample of Local Galaxies

The MASSIVE Survey samples the most massive K -band-selected early-type galaxies within the local 108 Mpc northern hemisphere (Ma et al. 2014). These galaxies have central stellar ages suggesting a formation epoch at $z > 2$ (Greene et al. 2015). Given the similar selection for our MQGs at $z > 2$, stellar masses, and inferred formation epoch, this sample is adopted as the local reference sample. This sample is further motivated in Section 5.1.

The extinction-corrected absolute K -band magnitudes listed in Table 3 of Ma et al. (2014) are converted into stellar masses using Equation (1) in van de Sande et al. (2019). The NASA-Sloan Atlas semimajor axis optical effective radii, also listed in Table 3 of Ma et al. (2014), are used. These were derived from 2D Sérsic (1968) fits with Sérsic parameters varying between $n = 2$ and 6. For the galaxies where this is not available, the infrared Two Micron All Sky Survey (2MASS) measurements were used to convert these to semimajor axis optical effective radii using Equation (4) in Ma et al. (2014). These sizes were derived from single Sérsic and de Vaucouleurs profile fits ($n = 4$). The effective velocity dispersion measurements used are reported in Veale et al. (2018). They were estimated using

²¹ https://docs.scipy.org/doc/scipy/reference/generated/scipy.stats.anderson_ksamp.html

Table 1
Summary of Sample

| Target ID | R.A. (deg) | Decl. (deg) | z_{phot} | Exp. Time | K | S/N_{HAB} | ESO Program | $(U - V)$ | $(V - J)$ |
|-------------------------|------------|-------------|-------------------|-----------|-------|-------------|---------------|-----------|-----------|
| UV-108899 | 150.17661 | 2.0608871 | 2.19 | 5.0 | 20.35 | 5.69 | 093.B-0627(A) | 1.60 | 0.80 |
| UV-250513 | 149.82227 | 2.6531196 | 2.03 | 5.0 | 20.37 | 4.12 | 093.B-0627(A) | 1.58 | 0.90 |
| CP-561356 | 150.20888 | 1.8502616 | 2.58 | 5.6 | 20.94 | 2.16 | 086.B-0955(A) | 1.63 | 0.82 |
| UV-105842 | 150.26265 | 2.0177791 | 1.93 | 4.0 | 20.20 | 4.28 | 093.B-0627(A) | 1.75 | 1.01 |
| UV-171687 | 149.88702 | 2.3506956 | 2.04 | 5.0 | 20.49 | 3.08 | 093.B-0627(A) | 1.37 | 0.94 |
| UV-90676 ^a | 150.48750 | 2.2700379 | 2.57 | 5.0 | 20.22 | 5.34 | 093.B-0627(A) | 1.53 | 0.81 |
| CP-1291751 | 149.86954 | 2.3167057 | 1.77 | 7.2 | 21.40 | 1.80 | 086.B-0955(A) | 2.19 | 1.19 |
| UV-155853 | 149.55630 | 2.1672480 | 1.96 | 5.0 | 20.36 | 4.65 | 093.B-0627(A) | 1.85 | 1.05 |
| UV-171060 ^b | 149.78951 | 2.3413286 | 2.02 | 5.0 | 20.45 | 3.89 | 093.B-0627(A) | 1.62 | 0.90 |
| UV-230929 | 150.20842 | 2.7721019 | 2.09 | 6.0 | 20.44 | 6.46 | 093.B-0627(A) | 1.48 | 0.68 |
| UV-239220 | 149.43275 | 2.5106428 | 2.00 | 4.5 | 20.40 | 2.86 | 093.B-0627(A) | 1.64 | 1.05 |
| UV-773654 | 150.74574 | 2.0104926 | 1.96 | 5.0 | 20.40 | 2.97 | 093.B-0627(A) | 1.81 | 1.04 |
| CP-1243752 ^c | 150.07394 | 2.2979755 | 1.98 | 4.5 | 20.07 | 5.25 | 086.B-0955(A) | 1.80 | 0.94 |
| CP-540713 | 150.32512 | 1.8185385 | 2.04 | 4.8 | 21.11 | 2.98 | 086.B-0955(A) | 1.61 | 0.82 |
| UDS 19627 ^d | 34.57125 | -5.3607778 | 2.02 | 5.0 | 20.19 | 4.40 | X-shooter GTO | 1.36 | 0.79 |

Notes. Target ID, right ascension (R.A.), declination (decl.), photometric redshift, X-shooter NIR arm exposure time in hours, total K magnitude, median S/N (9 \AA pixel^{-1} bins) in the H band ($15000 < \lambda[\text{\AA}] < 18000$), ESO program ID, rest-frame $(U - V)$ and $(V - J)$. The R.A., decl., photometric redshift, K band, and UVJ colors are from Muzzin et al. (2013a; except UDS 19627^d).

^a Previously published in Kado-Fong et al. (2017), Mowla et al. (2018), Marsan et al. (2019).

^b Previously published in Mowla et al. (2018).

^c Previously published in van de Sande et al. (2013), Krogager et al. (2014), Belli et al. (2014b, 2017), Allen et al. (2015), Kriek et al. (2016), Mowla et al. (2018).

^d Previously published in Toft et al. (2012; all values in table taken from there).

the MILES stellar library (Falc3n-Barroso et al. 2011) together with the penalized pixel-fitting algorithm (pPXF; Cappellari & Emsellem 2004). Finally, the average luminosity-weighted stellar velocity dispersion within the effective radius is adopted.

3. Data

Here we describe the spectroscopic observations with the VLT/X-shooter spectrograph (D’Odorico et al. 2006; Vernet et al. 2011) and the *HST*/WFC3 follow-up of our MQGs. These spectroscopic and photometric campaigns spanned an interval of more than 10 yr, spread over several programs that are summarized in Table 1. Finally, the ancillary data used in the analysis are presented.

3.1. VLT/X-shooter Spectroscopy

Mounted on the VLT, X-shooter is a single-object echelle spectrograph covering 3000–25000 \AA with three arms: UVB (2936–5930 \AA), VIS (5253–10489 \AA), and NIR (9827–24807 \AA). We were granted 35 and 57 service mode hours in P86 and P93, respectively (PI: Toft). The latter carried over and finished in period 96. The observations were completed using a default nodding mode to ensure a robust sky subtraction of the NIR band, probing the rest-frame optical part of the spectra for the $z \sim 2$ quiescent galaxies. The majority of the P86/P93 observations (89/96%) were completed with an average airmass-corrected DIMM seeing of $0''.8$ in the NIR arm. The telluric standard stars were observed close to the science observations, both in airmass and in time, to mimic the conditions of the sky and optimize the atmospheric absorption correction. The P86/P93 observations for the NIR (VIS) frames were executed with 480/900 s (314/863 s) exposures, a $0''.9 \times 11''$ slit configuration, and—for the P93 sample only—including the K -band blocking filter. We aligned the slit along the galaxy’s major axis in the UltraVISTA K -band images avoiding bright nearby sources.

The data are reduced using a wrapper of the ESO X-shooter pipeline (Modigliani et al. 2010; Sparre 2015), along with

customized modifications (Zabl et al. 2015). Beyond the standard pipeline processing steps for the NIR arm in nodding mode, we account for the spatial variations of the background level outside of the orders in each raw science frame by removing the median level obtained from the illuminated areas from each row of pixels in the detector. The 2D VIS and NIR individual science frames are corrected for telluric absorption with a customized and publicly available wrapper²² (Selsing et al. 2016) of the pPXF (Cappellari & Emsellem 2004), based on the PHOENIX stellar atmosphere library (Husser et al. 2013). A response function is constructed by modeling the atmosphere during the science exposures, and each individual observation block (OB) is corrected.

Finally, individual OBs are combined into an optimally weighted 2D spectrum removing flux outliers using 3σ and 5σ median clipping for the VIS and NIR, respectively. Bad pixels automatically flagged during the reduction are also excluded. Furthermore, off-trace emission is flagged and excluded in the construction of the OBs from UV-105842, UV-171687, and UV-155853 to minimize the contamination from surrounding sources. The 1D spectrum is optimally extracted (Horne 1986). Flux corrections are made, anchoring the synthetic photometry to the total magnitudes from the latest COSMOS15 catalog (Laigle et al. 2016, their Section 3.3) and accounting for point-spread function (PSF) matching in different bands and for the Galactic extinction. The H - and I -band magnitudes are used to compute independent aperture correction factors for the NIR and VIS spectra, respectively.

3.2. HST/WFC3 H_{F160W} Imaging

Eleven orbits of *HST*/WFC3 with *HST*-GO-14721 (PI: Conselice) are allocated to observe the rest-frame optical images, H_{F160W} , for UDS 19627 and the 10 galaxies in the P93

²² <https://github.com/jselsing/QuasarComposite/blob/master/py/telluric.py>

sample. The P86 sample is covered by the following programs: CP-1243752 (*HST*-GO-12440; PI: Faber) and CP-561356 (*HST*-HLA-14114; PI: van Dokkum) with WFC3 and CP-1291751 and CP-540713 with *HST*/NICMOS (*HST*-HLA-9999; PI: Scoville).

The WFC3/ H_{F160W} data are reduced using the “grism redshift and line” analysis software, Grizli,²³ which is an end-to-end processing code for WFC3/IR data using ASTRODRIZZLE.²⁴ The starting point is the standard calibrated images downloaded from the MAST archive (FLT extension images). The calibrated images are 1014×1014 pixels with $0''.13 \text{ pixel}^{-1}$. For each visit, there are four dithered exposures that are combined using Grizli. The resulting products for each visit are aligned, background-subtracted, and drizzled images with $0''.06 \text{ pixel}^{-1}$. The NICMOS data for CP-1291751 and CP-540713 are reduced in a similar manner with ASTRODRIZZLE.

3.3. Ancillary Data: Multiwavelength Photometry and HST I_{F814W} Images

We make ample use of the 14 broadband COSMOS photometry measurements from the Laigle et al. (2016) catalog, covering the full UV-to-NIR wavelength range to model our stellar populations in Section 4.3. The total magnitudes are adopted using the method described in Appendix A.2 by the same authors. Complementary to the UV-to-NIR photometry, we check the available deep X-ray *Chandra* imaging (Marchesi et al. 2016) and the “superdeblended” FIR catalog (Jin et al. 2018), superseding the previous $24 \mu\text{m}$ catalog (Le Floc’h et al. 2009) used in the selection of P93. This new implementation adopts active priors from the *Spitzer*/MIPS $24 \mu\text{m}$ and radio observations to deblend the low-resolution imaging from *Herschel*/PACS and SPIRE, SCUBA2, AzTEC, and MAMBO. The sources are cross-checked with the *Galaxy Evolution Explorer* far- and near-UV data from Zamojski et al. (2007) and Capak et al. (2007). This search for UV or X-ray counterparts results in no detections for any of our galaxies. On the other hand, we do find hints of mid-infrared (MIR) and radio emission from part of the sample, as detailed in Section 4.4.2 and discussed in Section 6.4. Similar UV-to-NIR multiwavelength coverage is available for UDS 19627. For an in-depth discussion of the available photometric data for this object, see Toft et al. (2012).

Thirteen of 15 galaxies have *HST* I_{F814W} imaging that is part of the COSMOS publicly released data (Scoville et al. 2007; Koekemoer et al. 2007). It covers $\sim 2 \text{ deg}^2$ of the sky with the Advanced Camera for Surveys (ACS) in the I band and comprises 81 tiles. Each tile is observed in four dithered exposures that are combined to produce a pixel scale of $0''.03 \text{ pixel}^{-1}$ and a PSF of $0''.095$ at FWHM. COSMOS images reach a point-source limiting depth of $AB(F814W) = 27.2 (5\sigma)$.

4. Analysis

We present in this section the analysis of our X-shooter spectra and *HST*/WFC3 H_{F160W} images. The spectroscopic redshift, velocity dispersion, and stellar population of our galaxies are measured by modeling the absorption features in the stellar continuum, together with the broadband photometry.

As we find no significant emission line detections in the spectra, we derive optical SFR upper limits (Section 4.4), which we compare with the estimates from the MIR photometry. The majority of the spatially offset sources caught in the spectra are foreground and background galaxies. Finally, the *HST* images probing the rest-frame optical structure are modeled to obtain their morphological parameters. The major merger candidates (UV-108899, UV-250513, and CP-561356; see Figure 2) are confirmed to be within redshift proximity such that their stellar masses can be reliably flux-corrected.

The *HST* red–blue (RB) color images, rest-frame optical X-shooter spectra with Laigle et al. (2016) photometry, and our best-fitting stellar population model are shown in Figure 2. For UDS 19627, the *HST*/WFC3 H_{F160W} image is presented in Section 4.5 and its spectrum is shown in Toft et al. (2012).

4.1. Spectroscopic Redshifts and Stellar Velocity Dispersion

All spectra of targeted sources (P86 and P93) show prominent hydrogen absorption features, which are typical of evolved stellar populations (see Figure 2). The stellar absorption features are modeled using pPXF, and both the line-of-sight velocity centroid (i.e., the spectroscopic redshift) and the line-of-sight stellar velocity dispersion (hereafter “velocity dispersion”) are measured.

The initial redshift and velocity dispersion guess is obtained from running pPXF with the Bruzual & Charlot (2003, hereafter BC03) stellar population library. The stellar population analysis is performed with complex SF history (SFH) fitting of the spectra and SED (see Section 4.3) adopting this initial estimate. The resulting best-fit model is confirmed to be stable against perturbations of $\Delta\sigma = \pm 100 \text{ km s}^{-1}$. The velocity dispersion measurement is refined by rerunning pPXF with a non-velocity-broadened best-fit stellar population model.

The spectra and best-fit model are convolved to the same resolution (FWHM = 3.2 \AA) and rebinned to a constant velocity scale without additional interpolation. Low-order additive ($a = 2$) and multiplicative ($m = 2$) correction polynomials are fit over the rest-frame range 3750–5950 \AA . The JH band gap and the regions where emission lines might be expected²⁵ are excluded while also masking out bad pixels.

The associated systematic and statistical errors are quantified by varying the wavelength range, correction polynomials, and stellar libraries (see details in Appendix B), similar to the method used in Toft et al. (2017). In all cases (P86 and P93), we determine secure redshifts, and for 10/14 galaxies, we estimate robust velocity dispersions. The spectroscopic redshifts and velocity dispersion measurements, along with the combined systematic and statistical errors (Appendix B), are listed in Tables 2 and 3, respectively. In Table 3, we also list the velocity dispersion for UDS 19627 derived in Toft et al. (2012). In Figure 3, the derived spectroscopic redshifts are compared with the photometric estimates from Muzzin et al. (2013a) and Laigle et al. (2016). Using the normalized median absolute deviation (σ_{NMAD}) from Brammer et al. (2008), no catastrophic outliers are found, except for photometric redshifts being systematically below the spectroscopic redshifts for both catalogs, finding a better agreement for Muzzin et al. (2013a).

²³ <https://github.com/gbrammer/grizli/>

²⁴ A Python implementation of Multidrizzle: <https://drizzlepac.readthedocs.io/en/latest/astrodizzle.html>.

²⁵ Excluded emission lines (wavelengths in \AA): [O II] (3726.03, 3728.82), [O III] (4958.92, 5006.84), [O I] (6300.30), [N II] (6548.03, 6583.41), H α (6563), and [S II] (6716.47, 6730.85).

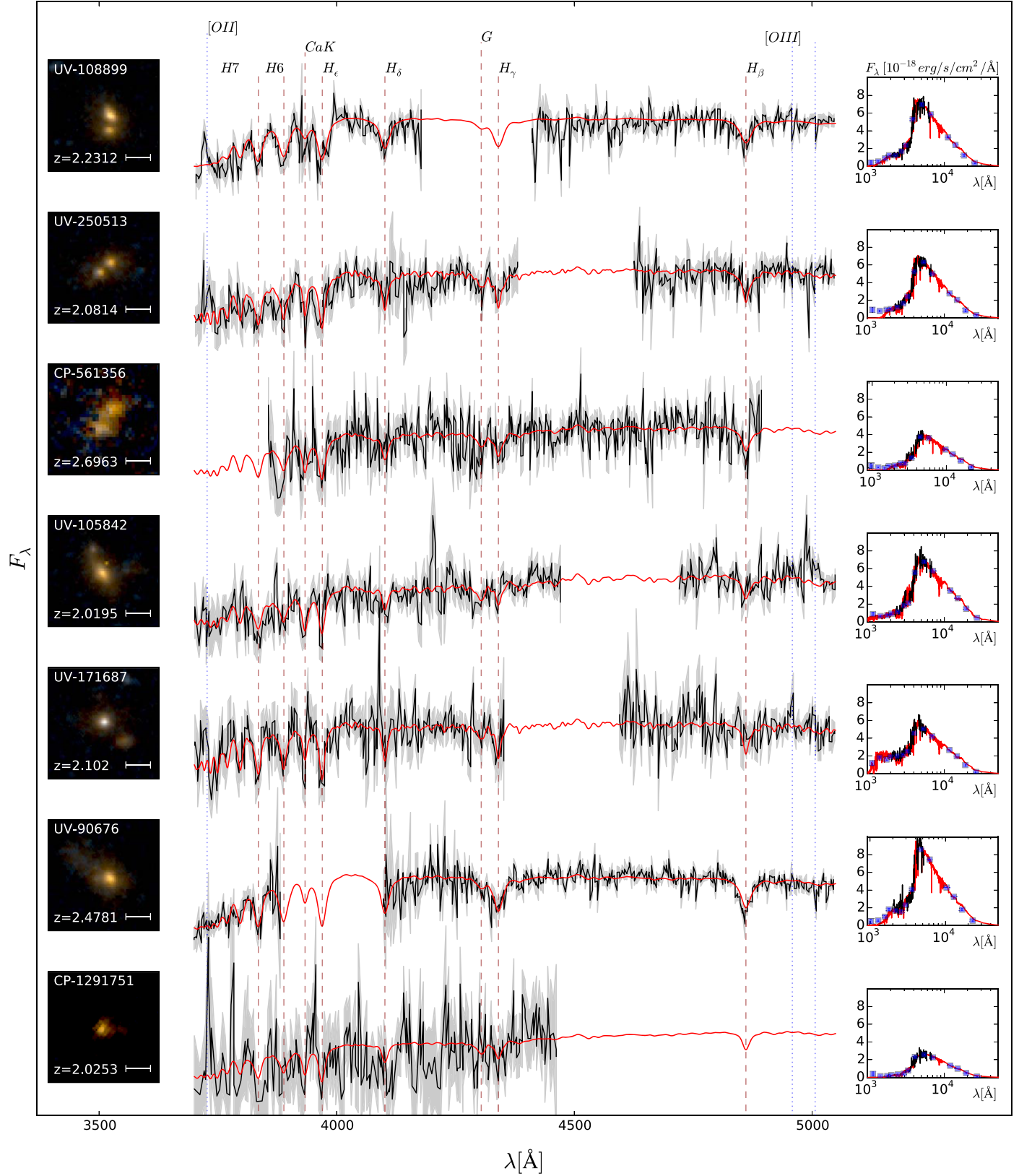


Figure 2. The VLT/X-shooter spectra of our sample, covering the rest-frame wavelength range $3700 \text{ \AA} < \lambda < 5050 \text{ \AA}$, with corresponding *HST* RB images (left column). On the right, the full SED is displayed by multiwavelength photometry (blue squares), and in the middle, the rest-frame optical X-shooter spectra (black line) and best-fit stellar population model (red line; Section 4.3) are given. Spectra are shown with an optimal adaptive binning and 1σ rms noise in gray shading. The two-color $4''.5 \times 4''.5$ northeast-orientated RB images, with galaxy ID and absorption line-determined spectroscopic redshifts (determined in Section 4.1), are made from *HST*/ACS I_{F814W} and WFC3 H_{F160W} . A $1''$ white bar is shown (~ 8.5 kpc at $z = 2$). The G, Ca K, and Balmer absorption features are indicated with dark red dashed lines, while [O II] 3727 Å and [O III] 4959, 5007 Å are indicated with blue dotted lines. The figure shows bright red sources with Balmer absorption lines, no significant optical emission lines, a strong 4000 Å break, and low rest-frame UV light, all indicative of quiescent stellar populations.

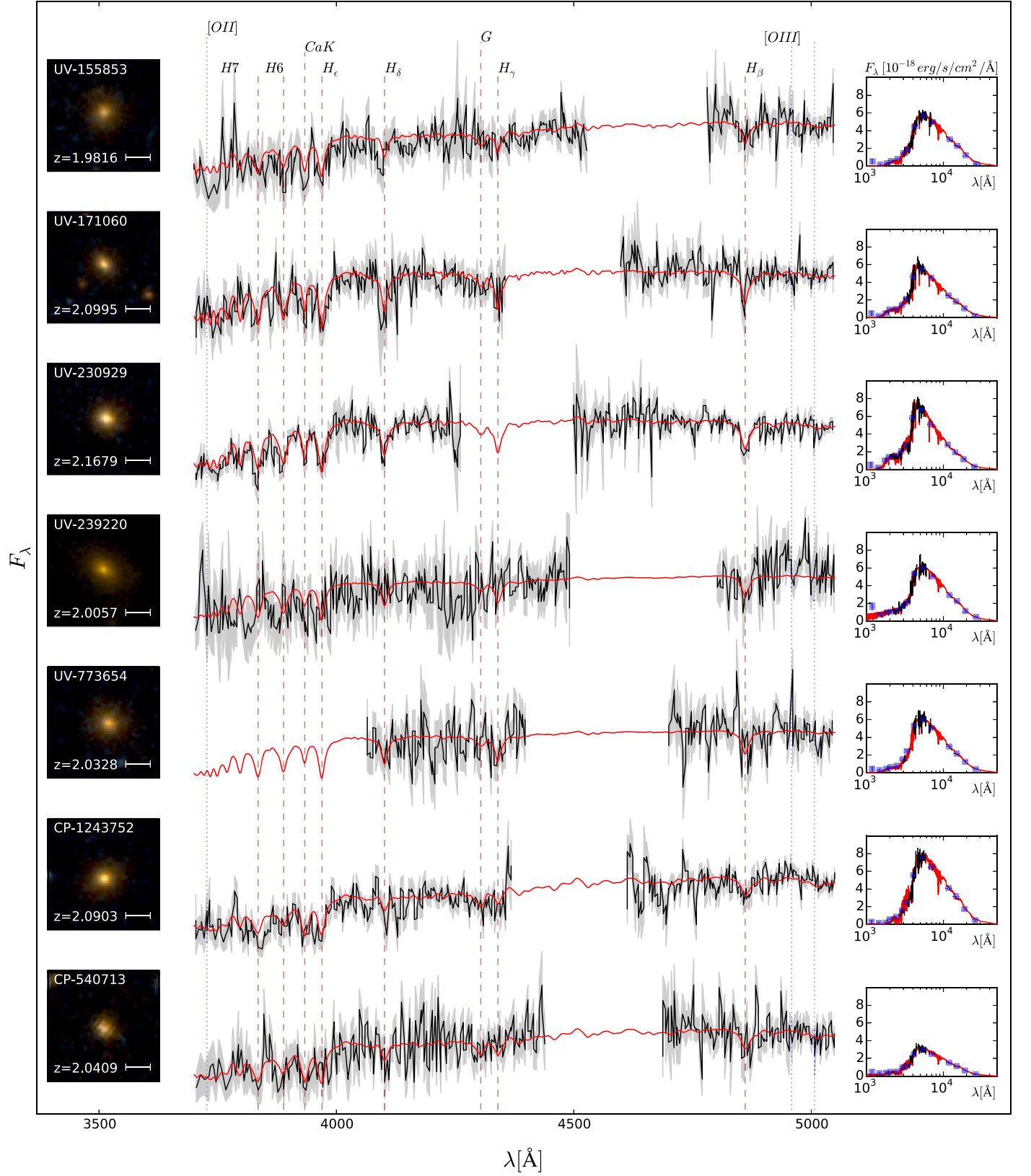


Figure 2. (Continued.)

4.2. Emission Lines

No on-source nebular line emission is detected at 3σ for any object in the sample. For UV-108899 and UV-239220, we find

indications of emission ($\sim 2\sigma$) from [O II] 3726.2, 3728.9 and H α 6563, respectively. In Appendix C, we discuss the specifics of the fitting method and list, in Table 2, the SFR and uncertainties from [O II] and H α (Kennicutt 1998).

Table 2
The Stellar Population Model Parameters

| Target ID (1) | z_{spec} (2) | $\log(M_*/M_\odot)$ (3) | $\log(\text{Age}/\text{yr})$ (4) | $A(g)$ (5) | $\text{SFR}_{\text{SSP}} (M_\odot \text{ yr}^{-1})$ (6) | $\text{SFR}_{\text{opt}} (M_\odot \text{ yr}^{-1})$ (7) | $\text{SFR}_{24} (M_\odot \text{ yr}^{-1})^a$ (8) |
|------------------------|--------------------------|----------------------------|-------------------------------------|------------------------|--|--|--|
| UV-108899 | 2.2312 | $11.62^{+0.16}_{-0.18}$ | $9.15^{+0.27}_{-0.30}$ | $0.38^{+1.00}_{-0.38}$ | <13 | 6 ± 4 ([O II]) | <15 |
| UV-250513 | 2.0814 | $11.51^{+0.18}_{-0.19}$ | $9.16^{+0.27}_{-0.31}$ | $0.38^{+1.02}_{-0.38}$ | <12 | <3 (H α) | <13 |
| CP-561356 | 2.6963 | $11.62^{+0.21}_{-0.20}$ | $9.14^{+0.28}_{-0.32}$ | $0.62^{+1.17}_{-0.62}$ | <86 | <19 ([O II]) | <90 [⊕] |
| UV-105842 | 2.0195 | $11.68^{+0.16}_{-0.17}$ | $9.19^{+0.26}_{-0.33}$ | $0.81^{+1.04}_{-0.81}$ | <17 | <2 (H α) | 19 ± 5 |
| UV-171687 | 2.1020 | $11.51^{+0.18}_{-0.19}$ | $9.13^{+0.28}_{-0.32}$ | $0.64^{+1.13}_{-0.64}$ | <24 | <3 (H α) | $26 \pm 6^{\diamond\oplus}$ |
| UV-90676 | 2.4781 | $11.78^{+0.17}_{-0.18}$ | $9.09^{+0.29}_{-0.29}$ | $0.41^{+0.99}_{-0.41}$ | <88 | <6 ([O II]) | <92 [⊕] |
| CP-1291751 | 2.0253 | $11.24^{+0.23}_{-0.22}$ | $9.17^{+0.27}_{-0.33}$ | $0.82^{+1.26}_{-0.82}$ | <17 | <2 (H α) | 18 ± 4 |
| UV-155853 | 1.9816 | $11.62^{+0.18}_{-0.17}$ | $9.23^{+0.24}_{-0.33}$ | $0.88^{+1.04}_{-0.86}$ | <14 | <4 (H α) | <15 |
| UV-171060 | 2.0995 | $11.48^{+0.16}_{-0.17}$ | $9.16^{+0.27}_{-0.31}$ | $0.41^{+1.03}_{-0.41}$ | <14 | <2 (H α) | <15 [⊕] |
| UV-230929 | 2.1679 | $11.48^{+0.16}_{-0.16}$ | $9.10^{+0.28}_{-0.28}$ | $0.22^{+0.89}_{-0.22}$ | <6 | <4 ([O II]) | <7 |
| UV-239220 | 2.0057 | $11.57^{+0.20}_{-0.20}$ | $9.18^{+0.26}_{-0.33}$ | $0.66^{+1.14}_{-0.66}$ | <19 | 35 ± 15 (H α) | $21 \pm 4^{\diamond\oplus}$ |
| UV-773654 | 2.0328 | $11.59^{+0.19}_{-0.20}$ | $9.20^{+0.26}_{-0.33}$ | $0.68^{+1.13}_{-0.68}$ | <12 | <2 (H α) | $13 \pm 3^{\diamond\oplus}$ |
| CP-1243752 | 2.0903 | $11.79^{+0.17}_{-0.17}$ | $9.23^{+0.24}_{-0.32}$ | $0.76^{+1.06}_{-0.76}$ | <11 | <2 (H α) | <12 |
| CP-540713 | 2.0409 | $11.26^{+0.22}_{-0.23}$ | $9.16^{+0.27}_{-0.32}$ | $0.57^{+1.19}_{-0.57}$ | <10 | <2 (H α) | <12 |
| UDS 19627 ^b | 2.0389 | $11.37^{+0.13}_{-0.10}$ | $9.08^{+0.11}_{-0.10}$ | $0.77^{+0.36}_{-0.32}$ | ... | <6 (H α) | <40 ^c |

Notes. Column 1: target ID. Column 2: spectroscopic redshift. Column 3: stellar mass. Column 4: mass-weighted stellar age. Column 5: extinction in g band. Column 6: 3σ upper limit percentiles (representing 99.73% Gaussian confidence intervals) of the stellar population modeled SFR/100 Myr distribution. Column 7: 3σ SFR upper limits based on H α or [O II] λ 3727 (Section 4.4.1). Column 8: 24 μm estimated SFR (see Section 4.4.2).

^a Galaxies with detections in 1.4 ($^\diamond$) and 3 ($^\oplus$) GHz are indicated with matching symbols.

^b The values listed for UDS 19627 are from Toft et al. (2012). From this study, the $A(v)$ extinction instead of the listed $A(g)$ is quoted.

^c The 2σ 24 μm SFR upper limit using the method from Franx et al. (2008).

Furthermore, spatially offset line emission is observed in four (UV-155853, UV-171687, UV-171060, UV-105842) 2D spectra coinciding with close-proximity sources. In 3/4 cases, this emission arises from foreground or background sources around UV-155853, UV-171687, and UV-171060 (Appendix A.2). The latter source northeast of UV-105842 shows significant [O II] 3726.2, 3728.9 Å, [O III] 4959, 5007 Å, and H α emission with a matching redshift of $z = 2.0124$. This corresponds to a velocity offset of $2130 \pm 120 \text{ km s}^{-1}$ from UV-105842. If purely due to galaxy motion, such an offset suggests that the two sources are not gravitationally bound at the time of observation. Another explanation of the asymmetric morphology might be a high-redshift analog of the locally observed offset AGN (Comerford & Greene 2014), likely caused by a recent merger event.

4.3. Stellar Population Modeling of Continuum Emission

In order to put constraints on the physical parameters of the stellar populations, the VIS+NIR X-shooter spectra and the broadband photometry are fit with the Bayesian approach from Gallazzi et al. (2005; recently revised in Zibetti et al. 2017) using the derived spectroscopic redshift. Spectral regions of poor atmospheric transitions are not included in the calculation. Before fitting, the models are convolved by the initial velocity dispersion estimated in Section 4.1.

Models are obtained by convolving the latest revision of the BC03 simple stellar population (SSP) models using the MILES stellar libraries (Sánchez-Blázquez et al. 2006; Falcón-Barroso et al. 2011) with a large Monte Carlo library of SFHs, metal-enrichment histories, and dust attenuations. The prior distribution of models is the one described in Zibetti et al. (2017), but here it is limited to 50,000 models with formation ages younger than 5 Gyr to be consistent with the high redshift of our galaxies. A full description of the model library is given in Zibetti et al. (2017); however, the most relevant information is summarized here.

The SFHs are modeled with a continuous component parameterized à la Sandage (1986),²⁶ thus allowing for both an increasing and a decreasing SFH phase, on top of which random bursts of SF are added. Stellar metallicity evolves according to the SFH (see Zibetti et al. 2019), with initial and final values randomly generated in the range 0.02–2.5 Z_\odot . Finally, for 75% of the models, the effect of dust attenuation is included following the model of Charlot & Fall (2000) that separates the contributions of the birth clouds affecting stars younger than 10^7 yr and the ISM affecting stars of all ages.

The Bayesian modeling approach assumes the likelihood of each model to be $\propto \exp(-\chi^2/2)$. The probability distribution function (PDF) of each physical parameter of interest is computed by weighing the prior distribution of the models in a given parameter by their likelihood, marginalizing over all other parameters. We additionally used the information from the MIR flux limit to restrict the sample of acceptable models to those that have an SFR consistent with 24 μm based upper limits and detections (see Section 4.4.2). The median and 16th and 84th percentiles of the PDFs are adopted as the fiducial estimates and their uncertainties for each parameter. Note that this approach allows the derivation of realistic uncertainties on the key physical parameters, accounting for both the observational errors and the intrinsic degeneracies among different parameters.

The stellar mass, mass-weighted mean stellar age, effective dust attenuation ($A(g)$), and SFR averaged over the last 100 Myr for our sample are reported in Table 2. In this table, the SFR limits from nebular line and 24 μm emission (see Sections 4.4.1 and 4.4.2) are also listed. Stellar masses are within the range of $\log_{10}(M_*/M_\odot) = 11.23 - 11.79$, with a median of 11.57. Compared to Belli et al. (2017), this sample is

²⁶ $\text{SFR}(t) = t/\tau \times \exp(-t^2/(2\tau^2))$; see, e.g., Section 3.1 in Zibetti et al. (2019).

Table 3
Summary of Structural Properties

| Target ID (1) | z_{phot} (2) | z_{spec} (3) | σ (km s $^{-1}$) (4) | R_{maj} (kpc) (5) | n (6) | q (7) | R_{flux}^a (8) | $\log_{10}(M_{*,c}/M_{\odot})$ (9) | $\log_{10}(M_{\text{dyn}}/M_{\odot})$ (10) | Class b (11) |
|--------------------------|--------------------------|--------------------------|---------------------------------|-------------------------------|------------|------------|----------------------------|---------------------------------------|---|--------------------|
| UV-108899-1 ^c | 2.19 | 2.2312 | 470 ± 82 | 1.36 ± 0.14 | 2.51 | 0.44 | 0.56 | $11.38^{+0.16}_{-0.18}$ | ... | P ^d |
| UV-108899-2 ^c | ... | ... | ... | 3.38 ± 0.34 | 7.15 | 0.56 | 0.44 | $11.26^{+0.17}_{-0.17}$ | ... | P ^d |
| UV-250513-1 ^c | 2.03 | 2.0814 | 174 ± 44 | 3.84 ± 0.38 | 4.00 | 0.59 | 0.55 | $11.26^{+0.17}_{-0.19}$ | ... | P ^d |
| UV-250513-2 ^c | ... | ... | ... | 1.60 ± 0.16 | 4.00 | 0.63 | 0.45 | $11.16^{+0.17}_{-0.18}$ | ... | P ^d |
| CP-561356-1 ^c | 2.43 | 2.6963 | 280 ± 128 | 4.14 ± 0.41 | 1.45 | 0.62 | 0.71 | $11.47^{+0.21}_{-0.20}$ | ... | P ^d |
| CP-561356-2 ^c | ... | ... | ... | 2.78 ± 0.28 | 0.90 | 0.39 | 0.29 | $11.09^{+0.21}_{-0.21}$ | ... | P ^d |
| UV-105842-1 | 1.93 | 2.0195 | 263 ± 57 | 4.07 ± 0.41 | 3.51 | 0.51 | 1.00 | $11.68^{+0.16}_{-0.17}$ | 11.61 ± 0.19 | P |
| UV-171687-1 | 2.04 | 2.1020 | 182 ± 50 | 5.12 ± 0.51 | 4.00 | 0.77 | 1.00 | $11.51^{+0.18}_{-0.19}$ | 11.37 ± 0.24 | P |
| UV-90676 | 2.57 | 2.4781 | 347 ± 82 | 5.22 ± 0.51 | 4.98 | 0.61 | 1.00 | $11.78^{+0.17}_{-0.18}$ | 11.89 ± 0.21 | P |
| CP-1291751 | 2.06 | 2.0253 | ... | 3.47 ± 0.35 | 3.59 | 0.67 | 1.00 | $11.24^{+0.23}_{-0.22}$ | ... | P |
| UV-155853 | 1.96 | 1.9816 | 247 ± 30 | 4.55 ± 0.46 | 3.62 | 0.85 | 1.00 | $11.62^{+0.18}_{-0.17}$ | 11.60 ± 0.11 | E |
| UV-171060 | 2.02 | 2.0995 | ... | 1.73 ± 0.17 | 4.00 | 0.54 | 1.00 | $11.48^{+0.16}_{-0.17}$ | ... | E |
| UV-230929 | 2.09 | 2.1679 | 252 ± 21 | 1.74 ± 0.17 | 3.01 | 0.73 | 1.00 | $11.48^{+0.16}_{-0.15}$ | 11.23 ± 0.08 | E |
| UV-239220 | 2.00 | 2.0057 | ... | 5.35 ± 0.54 | 4.21 | 0.62 | 1.00 | $11.57^{+0.20}_{-0.20}$ | ... | E |
| UV-773654 | 1.96 | 2.0328 | ... | 3.77 ± 0.38 | 3.34 | 0.84 | 1.00 | $11.59^{+0.19}_{-0.19}$ | ... | E |
| CP-1243752 | 2.01 | 2.0903 | 350 ± 53 | 2.85 ± 0.29 | 4.50 | 0.79 | 1.00 | $11.79^{+0.17}_{-0.17}$ | 11.66 ± 0.14 | E |
| CP-540713 | 1.98 | 2.0409 | 353 ± 97 | 1.65 ± 0.17 | 0.96 | 0.79 | 1.00 | $11.26^{+0.22}_{-0.23}$ | 11.59 ± 0.24 | E |
| UDS 19627 | 2.02 | 2.0389 | 318 ± 53 | 2.00 ± 0.20 | 3.32 | 0.51 | 1.00 | $11.37^{+0.13}_{-0.10}$ | 11.48 ± 0.15 | E |

Notes. Column 1: target ID. Column 2: photometric redshift from Muzzin et al. (2013a). Column 3: spectroscopic redshift (Section 4.1). Column 4: stellar velocity dispersion measurement (Section 4.1). Column 5: effective semimajor axis (Section 4.5). Column 6: Sérsic index (Sérsic 1968). Column 7: axis ratio $q = b/a$ (defined by the ratio between the semiminor and major axes). Column 8: flux scaling used to estimate the corrected stellar mass, $\log_{10} M_{*,c}$ (Section 4.6). Column 9: corrected stellar mass (Section 4.6). Column 10: dynamical mass (Section 5.4). Column 11: morphological classification (P: peculiar; E: elliptical) from Conselice et al. (2005).

^a Relative flux ratio = $F_i/(F_i + F_j)$.

^b Galaxies marked with (d) are classified as major mergers in Section 4.6.

^c Double sources have similar photometric and spectroscopic redshifts, as well as the stellar velocity dispersions estimated from their composite spectra.

^d Galaxies classified as major mergers in Section 4.6.

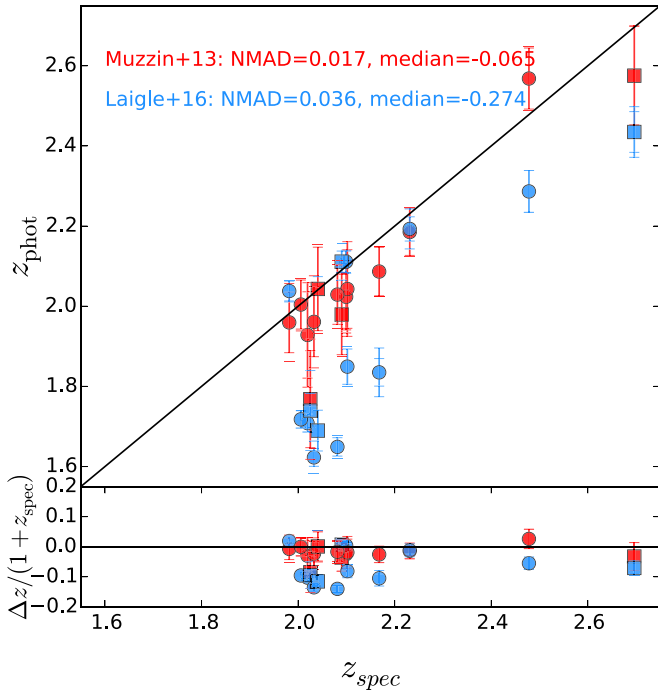


Figure 3. Comparison of spectroscopic and photometric redshifts for our sample of MQGs using the Muzzin et al. (2013a; red) and Laigle et al. (2016; blue) catalogs. The Muzzin et al. (2013a) catalog provides better photometric redshift estimates for MQGs at $z > 2$ compared to Laigle et al. (2016). Note that UDS 19627 is not in the same area of the sky covered by the catalogs compared here.

on average more massive, which is reflected by the brighter K -band magnitudes (see Figure 1). Such MQGs have also been found over a larger area in Arcila-Osejo et al. (2019). The SFR limits and dust-corrected stellar masses, together with the mean stellar mass-weighted ages of ~ 1.4 Gyr, confirm the expectations from the selection that this is, in fact, a sample of massive recently quenched galaxies. Three of the galaxies are double sources, and the stellar masses are corrected in Section 4.6.

4.4. SF and Quiescence

4.4.1. Rest-frame Optical Emission Lines

In order to confirm the quiescent nature of our galaxies, upper limits on [O II] $\lambda 3727$ and $H\alpha$ emission are measured. These are converted into upper limits of the unobscured SFRs following Equations (2) and (3) in Kennicutt (1998), under the assumptions of solar abundance ratio and that all massive SF is traced by ionized gas. A 3σ flux upper limit is determined by summing up the flux error density squared over a region of $\Delta\lambda = 1000$ km s $^{-1}$ (similar to 300–500 km s $^{-1}$ line dispersions):

$$F_{3\sigma \text{ limit}} = 3\sqrt{\sum \sigma_{\text{flux}}^2 \delta\lambda^2}. \quad (1)$$

Here σ_{flux} and $\delta\lambda$ are the flux uncertainty and bin size, respectively. Note that we do not introduce any dust extinction in this conversion, as this is largely unconstrained (see Section 4.4.3 for an estimated upper limit on the dust

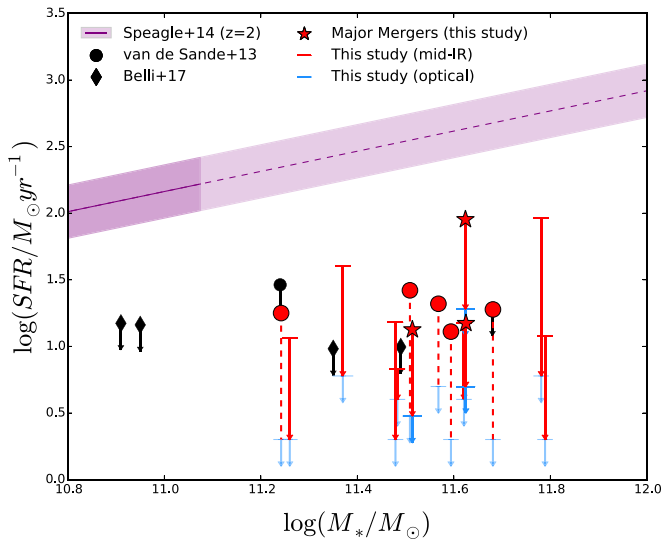


Figure 4. The SFR– M^* plane for MQGs at $z > 2$ with $24\ \mu\text{m}$ coverage. The SFR main sequence at $z = 2$ from Speagle et al. (2014) is shown in dark purple, with its $0.2\ \text{dex}$ (1σ) scatter. The light purple region extending beyond $\log(M_*/M_\odot) > 11.1$ is an extrapolation of the best-fit relation. The $24\ \mu\text{m}$ MIPS SFR detections (red circles)/upper limits (red arrows) are shown, with the major mergers (composite measurement of the SFR) shown by red stars. We show our rest-frame optical SFRs (based on $[\text{O II}]$ and $\text{H}\alpha$) in blue upper limits. We show $24\ \mu\text{m}$ SFR upper limits for the two objects from van de Sande et al. (2013; circles), together with four dust-corrected $\text{H}\alpha$ upper limits from Belli et al. (2018; diamonds) with black upper limits. Our sample of galaxies has suppressed SFR compared to the main sequence at $z = 2$ and can be considered truly quiescent galaxies.

extinction). We find unobscured SFR upper limits that are consistent with the expectation that these galaxies are quiescent ($-10 < \log_{10}(\text{sSFR}/\text{yr}) < -11.5$). The difference between the $[\text{O II}] \lambda 3727$ and $\text{H}\alpha$ SFR limits is $< 11\ M_\odot\ \text{yr}^{-1}$, and in Table 2, the lowest SFR upper limits are listed.

4.4.2. MIR Emission

The SFR, derived from rest-frame optical emission lines, represents a lower limit to the total SF in the presence of strong dust attenuation. Therefore, the SFR from the *Spitzer*/*MIPS* $24\ \mu\text{m}$ emission (Wu et al. 2005; Zhu et al. 2008; Kennicutt et al. 2009; Rieke et al. 2009) is estimated under the assumption of zero or subdominant AGN emission. Here the $24\ \mu\text{m}$ flux densities (or 3σ upper limits for sources undetected at $24\ \mu\text{m}$) from the most recent “superdeblended” FIR COSMOS catalog (Jin et al. 2018) are adopted. To derive SFR estimates, the $z = 2$ main-sequence SED template of Magdis et al. (2012) is rescaled to the measured $24\ \mu\text{m}$ flux densities (or the 3σ upper limits) of our targets. The emerging total infrared luminosity (L_{IR}) of the templates is converted to SFR through the L_{IR} –SFR relation of Kennicutt (1998), tuned to the adopted Chabrier IMF of this study. Detections corresponding to a median SFR $\sim 20\ M_\odot\ \text{yr}^{-1}$ are found for five of the galaxies that are undetected in the $24\ \mu\text{m}$ catalog (Le Floc’h et al. 2009). The remaining galaxies are not individually detected, and we thus fix them to their 3σ upper limit. Here UV-90676 and CP-561356 have upper limits of $\lesssim 90\ M_\odot\ \text{yr}^{-1}$. Both galaxies show strong merger signatures (see Section 4.6). The derived $24\ \mu\text{m}$ SFRs are listed in Table 2.

4.4.3. Comparison of Different SF Tracers

Figure 4 shows the position of the sample of MQGs in the $\log(\text{SFR})$ – $\log(M_*)$ main sequence at $z = 2$. For reference, the SFR main sequence at matching redshift from Speagle et al. (2014) is shown, extrapolated to the stellar mass range $\log_{10}(M_*/M_\odot) > 11.1$ covered by our galaxies.

The rest-frame optical SFR limits are systematically lower than the MIR estimates (both probing 10–100 Myr timescales). This suggests that the star-forming regions are either strongly obscured and/or dominated by AGN dust heating (Fumagalli et al. 2014). Under the assumption of no AGN contribution to the heating that produces the MIR emission (see also Section 6.4), the dust extinction is estimated by comparing the obscured and unobscured SFR estimates, resulting in a mean extinction of $A(V) < 1$ –2, consistent with our SED fit derived $A(g)$ (g -band) extinction. In order to judge if a significant contribution to the MIR heating arises from AGNs, we check whether there are any radio counterparts detected in Jin et al. (2018). Radio emission is detected in five sources at 1.4 GHz and five sources at 3 GHz (indicated with symbols in Table 2), showing that AGN heating could be responsible for the elevated MIR SFR estimates. Further treatment of the radio detections will be part of a future paper (I. Cortzen et al. 2020, in preparation).

The SFRs derived from our stellar population analysis (Section 4.3) are consistent with $\text{SFR} \sim 0\ M_\odot\ \text{yr}^{-1}$ for all galaxies in our sample. In Table 1, we list the 3σ upper limits on these SFR limits. However, even considering the most conservative upper limits on the SFR from the $24\ \mu\text{m}$ emission, our sample of MQGs lies $\sim 2\ \text{dex}$ below the SFR main sequence at their redshifts, confirming their quiescent nature (see Figure 4).

4.5. Galaxy Structure and Sizes

The 2D stellar light distribution traced by *HST*/WFC3 H_{F160W} imaging is modeled with the χ^2 -minimization fitting code GALFIT (Peng et al. 2002) in order to retrieve the structural parameters of our sample of MQGs. A first run of SExtractor (Bertin & Arnouts 1996) allows us to detect the objects in each field and obtain an initial guess for the structural parameters. A postage stamp for each target is constructed such that it encloses an ellipse with a major axis 2.5 times the Kron radius obtained by SExtractor. The local sky level in each stamp is calculated using Galapagos (Barden et al. 2012). This sky level is passed to GALFIT and kept fixed during the fitting. For the WFC3 data, a combination of the TINYTIM²⁷ simulated PSF and an empirical stacked star PSF are used. For the NICMOS data, an empirical stacked PSF is used.

Finally, GALFIT is run on each postage stamp, adopting a flexible Sérsic profile for every source (Sérsic 1968),

$$\Sigma(R) = \Sigma_e \exp \left\{ -\kappa_n \left[\left(\frac{R}{R_e} \right)^{1/n} - 1 \right] \right\}. \quad (2)$$

The parameter R_e is the effective radius enclosing half of the flux from the model light profile, $\Sigma(R_e)$ is the surface brightness at the effective radius, and n is the Sérsic index. The quantity κ_n is a function of the Sérsic index, which defines the global curvature of the light profile and is obtained by solving the equation $\Gamma(2n) = 2\gamma(2n, \kappa_n)$, where Γ and γ are,

²⁷ <http://www.stsci.edu/hst/observatory/focus/TinyTim>

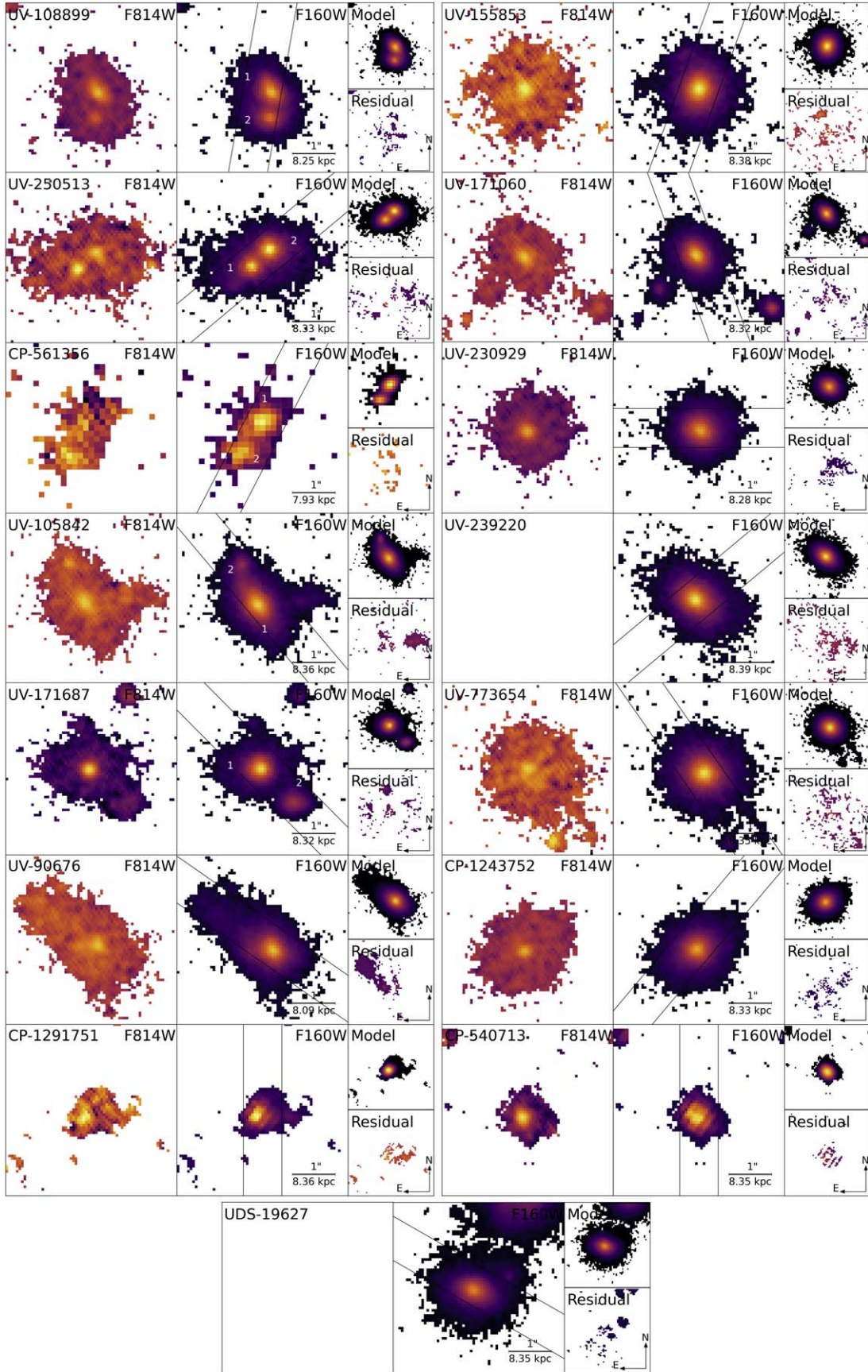


Figure 5. The I_{F814W} , H_{F160W} , GALFIT model, and GALFIT residual for our sample of MQGs at $z > 2$ in $4'' \times 4''$ cutouts. Pixels with 3σ confidence (with regard to background) are indicated with a logarithmic color scale to showcase the structure and morphology of the sample. The H_{F160W} significant pixels are used as a mask for all images. In the residual image, the pixels, one standard deviation above the background, are shown within this mask. The X-shooter slit is overlaid at the orientation of the spectroscopic observations. A scale of $1''$ is shown in kpc for size reference.

respectively, the gamma function and incomplete gamma function.

GALFIT is run several times to ensure that the solutions correspond to a global minimum in the minimization algorithm for each image by varying the initial guesses of the total magnitude, effective radius, and Sérsic index. The parameters are constrained so as to avoid any unphysical solutions (effective radius > 0.2 pixels, $q > 0.1$, $0.5 < n < 8$). Initially, all targets are fit with n as a free parameter. In unstable cases where the maximum or minimum n is reached, the images fixing the Sérsic index at either $n = 1$ or $n = 4$ are refit, choosing the model providing the smallest χ^2 as the best-fit solution. These two choices represent realistic descriptions of an early-type galaxy dominated by either a disk or a bulge. Throughout the whole fitting procedure, neighboring objects are either modeled or masked, depending on their proximity to the main target. A 10% measurement uncertainty on the size (van der Wel et al. 2008; Newman et al. 2012) is shown to be a fair representation. This conservative error estimate is thus adopted. The semimajor axis, $R_{e,\text{maj}}$, is adopted as the effective radius in the following sections. The best-fit parameters and their uncertainties are reported in Table 3.

In Figure 5, we present the rest-frame UV (I_{F814W}) and optical (H_{F160W}) images along with the GALFIT model and residual. The morphologies of these galaxies are classified in the H_{F160W} image according to Conselice et al. (2005), and they fall into the two categories for quiescent systems: ellipticals (E) and peculiars (P). When available, the spectroscopic observations are used to determine the distance in redshift space to objects that fall in the X-shooter slit (see Section 4.2). The majority of sources turn out not to be associated with the central galaxy. Nine of 15 galaxies are categorized as elliptical galaxies, while the remaining ones are categorized as peculiar galaxies with major mergers (UV-108899, UV-250513, CP-561356), minor mergers (UV-105842, CP-1291751), and/or strong tidal/post-merger features (UV-105842, UV-90676). The galaxies UV-108899, UV-250513, and CP-561356 are confirmed as ongoing major mergers in the following section. The classifications and morphological parameters are listed in Table 3.

4.6. Spectroscopic Confirmation and Stellar Mass Correction of Ongoing Major Mergers

The RB color images in Figure 2 reveal that three galaxies (UV-108899, UV-250513, CP-561356) appear to be double systems. The spectra, shown in the same figure, are the total extraction of the combined light from the two galaxies. These objects are within close proximity, and the light in the reduced 2D frames is blended to an unknown extent (due to limited seeing). At the expense of drastically decreasing the signal-to-noise ratio (S/N), an attempt to separate the sources and determine whether their individual redshift measurements can confirm their proximity is made.

For each system, the resolved 1D *HST* H_{F160W} light profile (extracted parallel to the X-shooter slit) is overlaid on the wavelength-collapsed 2D spectrum trace. A double Gaussian profile fit allowed us to gauge the amount of blending and make a conservative extraction of each individual galaxy, minimizing cross-source contamination. In Figure 6, the individual extractions and the best fit to the composite spectrum from Section 4.3 are shown.

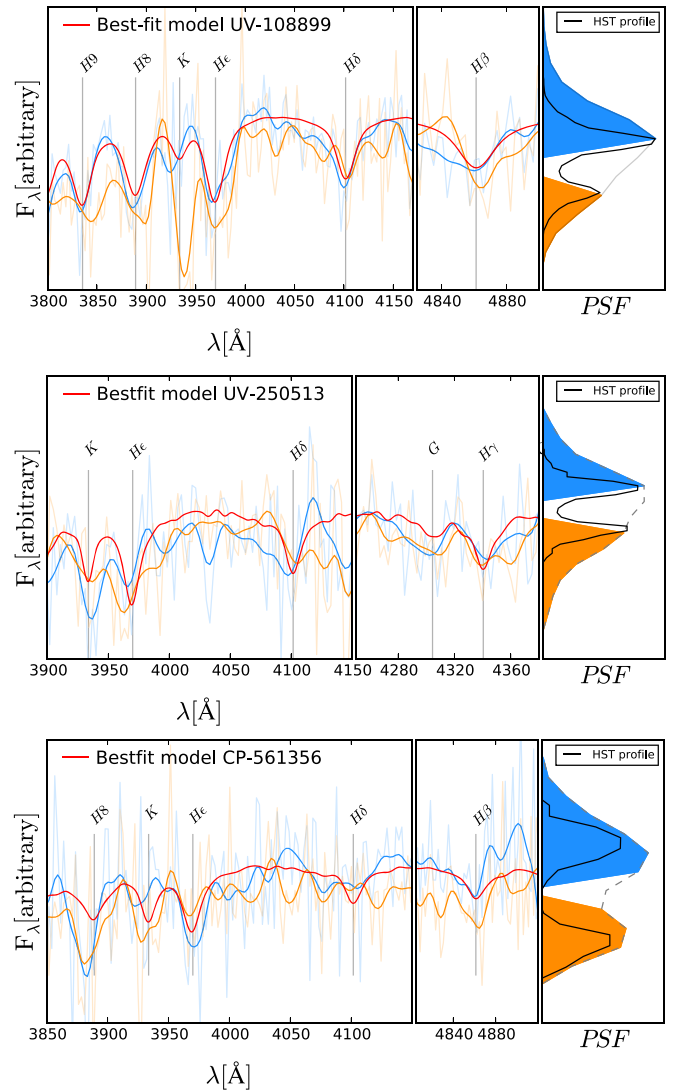


Figure 6. Individual flux extractions (blue, orange) from spatially divided 2D spectra of the major merger candidate sources UV-108899, UV-250513, and CP-561356 (top to bottom). The thick line is a smoothed version of the original spectra shown by the thin line. The right panel shows the wavelength-collapsed 2D spectrum (gray line) color-coded to match the individual extracted 1D spectra (left and middle panels). For reference, the 1D resolved *HST* H_{F160W} profile is shown by a thin black line. The best-fit model of the composite spectrum is shown in red, and the visible Balmer absorption lines are indicated. For each galaxy, we confirm the spectroscopic redshift proximity by the matching of absorption lines and conclude that these sources are ongoing major mergers.

Because of the low S/N of the individual conservative flux extractions, the estimation of the velocity offset is refrained, since it would be dominated by large uncertainties. However, the galaxies are within close physical proximity due to the matching absorption lines shown in the figure and can be considered ongoing quiescent (dry) major mergers. This confirmation is important as, in the following section, it can be used to correct their stellar masses, prior to presenting them in the mass–size plane (see Section 5.3).

Spectroscopic confirmation allows us to deblend the composite stellar mass of each system using the H_{F160W} magnitude as a proxy for tracing the bulk of the stars in the galaxies. The GALFIT-modeled H_{F160W} flux ratio supports the

fact that these galaxies are major mergers with mass ratios of 1:1–3. We used the flux ratio to correct the stellar masses as

$$M_{*,i} = M_{*,\text{tot}} \frac{F_i}{F_i + F_j} = M_{*,\text{tot}} R_{\text{Flux}}, \quad (3)$$

where i and j refer to the two merging galaxies and F is the total flux from GALFIT. The corrected stellar masses ($M_{*,c}$) and the relative flux ratio scaling, R_{Flux} , are listed in Table 3, with sources names matching the numbering in Figure 5. Following this correction, the galaxies are still classified as MQGs with stellar masses, $\log_{10}(M_*/M_\odot) > 11$.

5. Results

5.1. Minimal Progenitor Bias

A major issue preventing us from deriving a consistent evolutionary picture connecting galaxy populations across time is the “progenitor bias” problem (e.g., van Dokkum & Franx 1996; Carollo et al. 2013). When comparing galaxies across time, the implicit assumption is that the high-redshift sample contains all progenitors of the low-redshift reference sample. However, the fraction of quenched galaxies has been found to grow over time (Buitrago et al. 2013), introducing an unknown bias when comparing samples of galaxies across different epochs.

One approach that has been suggested to minimize the progenitor bias is comparing the evolution of galaxies at fixed velocity dispersions (see, e.g., Belli et al. 2014a). Archaeological studies (Graves et al. 2009; van der Wel et al. 2009; Bezanson et al. 2012) find evidence suggesting that the velocity dispersion in quiescent galaxies remains approximately unchanged across cosmic time ($z < 1.5$). In such a scenario, the velocity dispersion must be weakly affected by the average merger history, which, according to the numerical study by Hilz et al. (2012), occurs for minor merger-driven evolution. A detailed discussion on fixed velocity dispersion evolution is given in Belli et al. (2014a, 2017). Another way to minimize the progenitor bias has been to study galaxy populations at a constant cumulative number density (CND) instead of a fixed velocity dispersion or stellar mass (see, e.g., Mundy et al. 2015). This approach is introduced in van Dokkum et al. (2010) and refined further in Behroozi et al. (2013) and Leja et al. (2013). In Section 2.1, a sample of massive galaxies with central stellar population ages suggesting formation at $z > 2$ is introduced. This sample is volume-limited and represents the most massive early-type systems observed in the local universe. In order to draw a meaningful comparison, a subgroup of the most massive galaxies at $z = 0$ is selected and matched with the CND at $z = 2$. This will hereafter be referred to as the “fixed” CND. This approach is based on the assumption that the rank of galaxies within the stellar mass function is not strongly affected across cosmic time. This occurs if the stellar mass continuously grows from $z = 2$ to 0, implying the availability of surrounding material to accrete (or events that trigger secondary SF, although this is not expected for the MQGs at $z > 2$; Brammer et al. 2011; Behroozi et al. 2013; Muzzin et al. 2013b; Marchesini et al. 2014).

First, the CND of massive ($\log(M_*/M_\odot) > 11.2$) UVJ quiescent galaxies in the redshift range $1.9 < z < 2.5$ is estimated using the Muzzin et al. (2013a) catalog. The stellar mass limit represents the lower limit on the standard deviation of the mean stellar mass from the sample of galaxies studied in

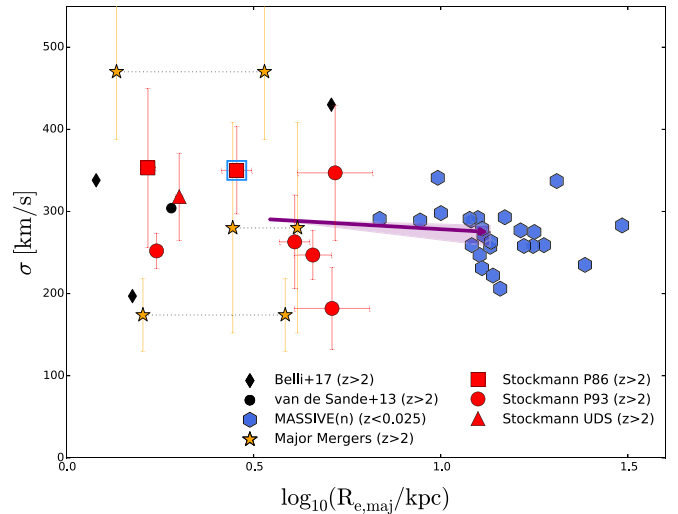


Figure 7. The velocity dispersions are plotted with effective radii for three samples: (1) our sample (red symbols, as in Figure 1); (2) massive, $\log_{10}(M_*/M_\odot) > 11$, other quiescent galaxies at $z_{\text{spec}} > 2$ (van de Sande et al. 2013; Belli et al. 2017); and (3) the MASSIVE(n) sample (blue hexagons). The composite dispersion measurements of the major merger galaxies are shown by orange stars with their individual size measurements connected by a horizontal dotted line. The blue square indicates our source, CP-1243752 (recently published in van de Sande et al. 2013; Kriek et al. 2016; Belli et al. 2017). The purple arrow shows the median evolution between our study and the MASSIVE(n) sample. The uncertainty from the mass rank scatter on the fixed CND is shown by purple shading. The median evolution between our study and the MASSIVE(n) sample shows evidence for shallow or no kinematic evolution from $z = 2$ to 0.

this paper. Our sample is 22% stellar mass complete using these selection criteria. We count 58 galaxies inside a comoving volume spanned by this redshift range, giving $n(\log(M_*/M_\odot) > 11.2) = 9.7 \times 10^{-6} \text{ Mpc}^{-3}$.

The MASSIVE galaxy sample is trimmed starting from the most massive object of the survey and including progressively less massive systems until we reach the fixed CND of the massive UVJ quiescent galaxies at $z \sim 2$. The final fixed CND-matched MASSIVE sample consists of the 25 most massive local elliptical galaxies with stellar masses of $\log(M_*/M_\odot) > 11.70$. The fixed CND-matched MASSIVE sample is hereafter referred to as “MASSIVE(n).” The MASSIVE(n) sample is considered a minimal progenitor-biased sample and used as our local reference sample in Sections 5.2–5.4.

The CND evolution suffers from large uncertainties from individual merger histories causing scatter in the mass rank, which is the main uncertainty for the highest stellar masses (Behroozi et al. 2013; Torrey et al. 2017). In Torrey et al. (2017), they estimated the mass rank scatter for $\log_{10}(M_*/M_\odot) > 11$ galaxies in Illustris (Genel et al. 2014; Nelson et al. 2015) by forward modeling of the CND. Their forward modeling, referred to as density distribution functions, is well described by a lognormal distribution, and the uncertainties can thus be treated as confidence intervals. For massive galaxies, the dominating uncertainty, the mass rank scatter, introduces a uncertainty of a factor of ~ 2 (within 80% confidence intervals) on the CND following the evolution from $z = 2$ to 0. In Behroozi et al. (2013), they found a similar uncertainty for the fixed CND evolution. This uncertainty on the CND evolution from the mass rank scatter is adopted and used to repeat the selection of the local reference sample, resulting in a corresponding uncertainty on the limit of the

stellar mass cut $\log(M_*/M_\odot) > 11.70^{+0.07}_{-0.10}$ and thus the number of galaxies in the local reference sample.

As an alternative approach to the fixed CND matching, the probabilistic approach from Wellons & Torrey (2017) is used to estimate the CND at $z = 0$. In Appendix D, the results (from Figures 7–9) for both a fixed and a probabilistic CND-matching approach is presented. The choice of CND-matching method does not affect the qualitative results of this paper presented in the following sections.

5.2. Kinematic Evolution of MQGs from $z = 2$ to 0

In Figure 7, the stellar velocity dispersion–size plane that allows us to study the kinematic evolution of MQGs from $z = 2$ to 0 is presented. The ongoing major merger galaxies are included to show that their incorrect composite dispersion measurements increase the scatter if not properly accounted for.

The mean velocity dispersion of the sample studied in this paper is $289 \pm 58 \text{ km s}^{-1}$ (without major mergers). This is consistent with previous $z > 2$ MQG literature (see studies shown in Figure 7) with a mean dispersion of $272 \pm 31 \text{ km s}^{-1}$. Our velocity dispersion and size measurements (including other structural parameters) for CP-1242752 (indicated by the blue square in Figure 7) are consistent with previously published values (van de Sande et al. 2013; Belli et al. 2014b, 2017; Kriek et al. 2016).

Comparing the median dispersion of our study to that of the local MASSIVE(n) sample, a shallow or no kinematic evolution from $z = 2$ to 0 is found. In Figure 7, significant effective size evolution consistent with earlier findings is observed (Newman et al. 2012; van der Wel et al. 2014). The effect of the mass rank scatter on the fixed CND matching is shown as the purple shading around the median evolution. This shading outlines the variation on the median when using the upper and lower limits of the CND matching (based on the stellar mass cut $\log(M_*/M_\odot) > 11.70^{+0.07}_{-0.10}$) from the mass rank scatter.

Half of the morphologies of compact massive galaxies at $z \sim 2$ have been suggested to be disk-dominated (van der Wel et al. 2011). So far, only one spatially resolved study of a rotating disk quiescent galaxy at this epoch has been discovered (Geier et al. 2013; Toft et al. 2017; Newman et al. 2018). The line-of-sight measured velocity broadening of the absorption lines could be a combination of both rotation and dispersion in the presence of a disk-dominated system (see an analytical prescription in Belli et al. 2017). Care must therefore be taken when comparing $z > 2$ spatially unresolved dispersion with resolved local measurements.

Wuyts et al. (2011) showed that the stellar light distribution of galaxies, measured by the Sérsic index, traces well the $\log(\text{SFR})$ – $\log(M_*)$ relation, separating disk and spheroidal galaxies by $n = 2.5$ at $z < 1.5$. Under the assumption that this is valid at $z = 2$, we classify our galaxies by Sérsic index and find that 92% of our galaxies have spheroidal ($n > 2.5$) morphologies (when excluding the ongoing major mergers). If the Sérsic index $n > 2.5$ is a good tracer of dispersion-dominated systems at $z > 2$, it suggests that our sample of galaxy dispersion measurements is not strongly contaminated by rotation. Bezanson et al. (2018) find no strong correlation between Sérsic index and rotational support. Instead they find that galaxies with stellar masses $\log(M_*/M_\odot) > 11.2$ (similar to our systems) appear to not be rotationally supported.

A recent study by Veale et al. (2018) presents the spatially resolved velocity dispersion measurements for the MASSIVE

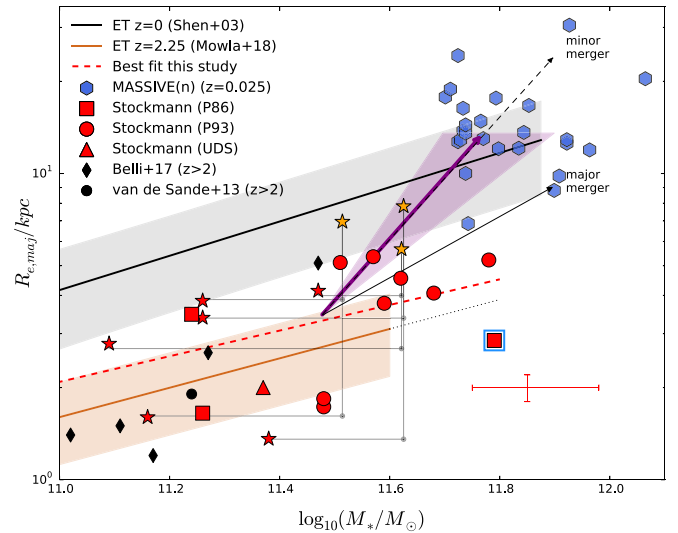


Figure 8. Stellar mass–size plane for massive, $\log_{10}(M_*/M_\odot) > 11.0$, quiescent galaxies: our sample (red symbols), other MQGs at $z > 2$ (black symbols; van de Sande et al. 2013; Belli et al. 2017), and the MASSIVE(n) sample (blue hexagons). The representative error bar of our sample is shown in red. The source CP-1242752 is indicated with a blue square. The ongoing major merger–corrected stellar masses (red stars) are connected (gray dotted, dashed, and solid lines) to their post-merger positions (orange stars) following the Bezanson et al. (2009) prescription. The minor (dashed) and major (solid) merger–predicted evolutions from Bezanson et al. (2009) are shown with black arrows. The best-fit relations at $z = 0$ (Shen et al. 2003) and 2.25 (Mowla et al. 2018), with their 1σ uncertainty, are shown in black and brown, respectively. The best-fit relation to the galaxies of this study is shown in dashed red. The purple arrow shows the median evolution between our study and the MASSIVE(n) sample. The shaded purple area represents the uncertainty on the median of the MASSIVE(n) sample when the mass rank scatter from Behroozi et al. (2013) is taken into account (see explanation in Section 5.2). The median mass–size evolution of MQGs from $z = 2$ to 0 can be explained primarily by minor mergers.

Survey sample. Here $\log_{10}(M_*/M_\odot) > 11.7$ galaxies (similar to our stellar mass cut of the MASSIVE(n) sample) all have velocity dispersions in the range $200 \text{ km s}^{-1} < \sigma < 350 \text{ km s}^{-1}$ at all radii (< 15 – 30 kpc). This rules out the possibility that the shallow dispersion evolution comparison is driven by spatial resolution. A comparison to the fixed CND-matched MASSIVE(n) sample establishes that the dispersion remains nearly unchanged.

A negligible median dispersion evolution of our MQGs across the last 10 billion yr ($z = 2$ – 0) is found in Figure 7 when comparing to the MASSIVE(n) sample. In the absence of spatially resolved spectroscopy, we make use of the morphological classification and high stellar masses, which suggests that our kinematics are unlikely to be strongly contaminated by rotation. Studying the evolution of galaxies at fixed dispersion has been suggested as a method to minimize progenitor bias (e.g., Belli et al. 2014b).

5.3. Stellar Mass–Size Plane for MQGs

In Figure 8, the stellar mass–size plane ($\log_{10} M_*$ – $R_{e,\text{maj}}$) is presented, which allows us to study the structural and stellar mass evolution of MQGs since $z \sim 2$. The three ongoing major merger galaxies with resolved sizes of individual galaxies (Section 4.5) and their flux-corrected stellar masses (Section 4.6) are shown in the figure. The post-merger stellar masses and sizes of these are predicted using the argument of virialization from Bezanson et al. (2009). The resulting position of post-merger galaxies is consistent with the average locus of

the most massive ($\log_{10}(M_*/M_\odot) > 11.5$) individual galaxies in our sample, showing that a way to form the most MQGs in our sample could be major quiescent-to-quiescent dry galaxy mergers (Naab et al. 2006).

A best-fit relation to the galaxies in this study, including the major merger-separated galaxies, reveals a shallower slope than what is found in the van der Wel et al. (2014) $z = 2.25$ mass-size relation but in better agreement with Mowla et al. (2018). The best-fit parameters, using a similar parameterization ($r/\text{kpc} = A(M_*/(5 \times 10^{10})^\alpha)$), are $\log(A) = 0.19$ and $\alpha = 0.42$. The stellar mass for CP-1243752 (blue square in Figure 8) is consistent within a 1σ standard deviation with van de Sande et al. (2013) and Belli et al. (2017) and within 1.1σ for the stellar mass published in Kriek et al. (2016).

The distribution of our sample shows that $z > 2$ MQGs are \sim two times more compact than objects with the same stellar mass in the local universe (Shen et al. 2003), which is a well-established result in previous works (van de Sande et al. 2013; Belli et al. 2017). The median stellar mass and size for our (MASSIVE(n)) sample, $\log(M_*/M_\odot) = 11.48$ (11.77) and $R_{e,\text{maj}}/\text{kpc} = 3.42$ (13.55), show that a doubling (~ 0.3 dex) in stellar mass and a factor of 4 in size evolution are required to bring the two samples into qualitative agreement.

Using the method from Bezanson et al. (2009) for predicting stellar mass and size growth, minor and major merger tracks are shown in the mass-size plane. The median mass-size evolution between our $z > 2$ MQGs and the local MASSIVE(n) sample could be explained by minor merger-predicted size and stellar mass growth. The tracks start at the median size and stellar mass of our sample (only red symbols). The qualitative conclusions remain the same when using a mean instead of a median or changing the choice of reference (with/without the major merger galaxies).

The median logarithmic mass-size slope is $\alpha = 1.78^{+0.37}_{-0.29}$ ($r \propto M_*^\alpha$). The uncertainties are determined based on the CND mass rank scatter, shown as the purple shaded area in Figure 8. This confirms the suggestion that minor mergers ($\alpha = 2$), compared to major mergers ($\alpha = 1$), are the preferred evolutionary path in the mass-size plane.

In line with earlier studies (van de Sande et al. 2013; van der Wel et al. 2014; Belli et al. 2017; Mowla et al. 2018), we find that our sample of $z > 2$ MQGs is compact in the stellar mass-size plane and further suggests that minor merger-driven size evolution (Bluck et al. 2012; Hilz et al. 2012, 2013; Newman et al. 2012; Oogi & Habe 2013; Fagioli et al. 2016) is preferred when comparing to the fixed CND-matched MASSIVE(n) sample.

5.4. Stellar-Dynamical Mass Plane for MQGs

In Figure 9, the dynamical-to-stellar mass relation for MQGs is plotted in order to study the interplay between the stellar and total (dynamical) mass potential over time. The dynamical mass derived from the Jeans equation (Jeans 1902) for symmetrical systems is as follows:

$$M(r) = \beta \frac{R_{e,\text{maj}} \sigma^2}{G}. \quad (4)$$

Here $R_{e,\text{maj}}$ is the effective semimajor axis, σ is the stellar velocity dispersion, G is the gravitational constant, and β is a parameter incorporating the full complexity of collisionless systems with radial-dependent parameters of density, dispersion, and velocity anisotropy. Following Cappellari et al. (2006), $\beta(n) = 8.87 - 0.831n + 0.0241n^2$ is adopted, where n is

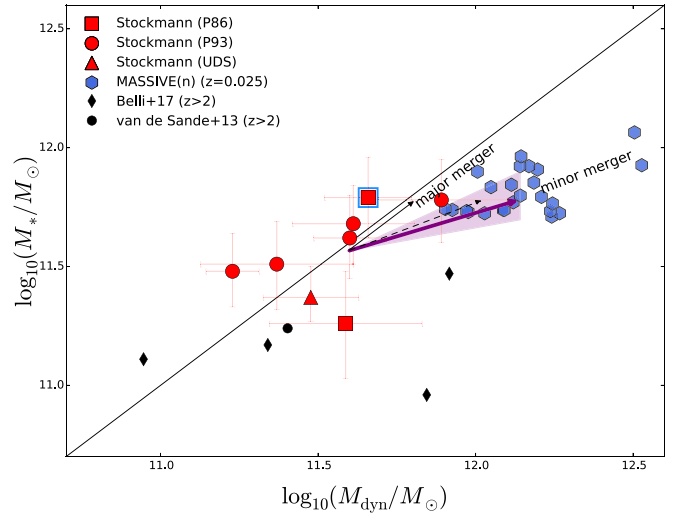


Figure 9. Dynamical-stellar mass plane for this study (red squares, P86; circles, P93; triangle, UDS 19627), other $z_{\text{spec}} > 2$ MQGs (black symbols; van de Sande et al. 2013; Belli et al. 2017), and the MASSIVE(n) sample (blue hexagons). The purple arrow connects the median of our sample with the median of the MASSIVE(n) sample. The purple shaded area represents the uncertainty on the median values of the MASSIVE(n) sample from the CND mass rank scatter (see explanation in Section 5.2). The solid black line is the $M_* = M_{\text{dyn}}$ relation. The dashed/solid black arrow represents the predicted constant dispersion stellar-to-dynamical mass evolution for minor/major mergers (Bezanson et al. 2009). The blue square indicates the source CP-1243752 (previously published in van de Sande et al. 2013; Kriek et al. 2016; Belli et al. 2017). The calculated dynamical-to-stellar mass ratio doubles from $z = 2$ to 0 when comparing to the fixed CND-matched MASSIVE(n) sample.

the Sérsic index (Sérsic 1968). The representation of β is a good approximation for symmetric systems, such as an elliptical galaxy that is well represented by a de Vaucouleurs profile. Taylor et al. (2010a) and Cappellari et al. (2013) showed that using such a parameterization of β yields dynamical masses in better agreement with the stellar masses when the sizes are estimated using a 2D Sérsic fitting method, rather than a fixed value of β .

The galaxies of this study are consistent with the stellar-to-dynamical mass ratio, $M_*/M_{\text{dyn}} < 1$, within the large uncertainties. A ratio > 1 is referred to as a nonphysical (forbidden) region, where the total mass is smaller than the mass of the stars. The galaxy, UV-230929, is located in this region at 1.1σ standard deviation from the $M_*/M_{\text{dyn}} = 1$ relation. Unfortunately, our large uncertainties prohibit trustworthy estimates of the total dust+gas mass for our sample. In Belli et al. (2017), it is suggested that dispersion-dominated systems with $n > 2.5$ lie closer to the $M_*/M_{\text{dyn}} = 1$ relation at $z \sim 2$.

Compared to previous $z > 2$ MQG studies (see legend in Figure 9), our sample occupies a similar dynamical mass range but has larger stellar masses. This is further discussed in Section 6.3. The dynamical mass for CP-1243752 (indicated by a blue square) is consistent with the previous measurements in van de Sande et al. (2013) and Belli et al. (2017).

A comparison between our study and the MASSIVE(n) sample is made to learn about the fixed CND evolution in the dynamical-stellar mass plane. The median evolution in Figure 9 illustrates that the dynamical mass evolves $2\times$ faster than the stellar mass within the effective radii. This means that the galaxies evolve such that the M_*/M_{dyn} ratio decreases from $z = 2$ to 0.

The minor and major merger evolution are shown for constant velocity dispersion evolution ($\Delta r \propto M_*^\alpha$), with $\alpha = 1$

for major merger and 2 for minor merger evolution. This is motivated by the shallow/constant dispersion evolution found in Section 5.2 when also comparing to the MASSIVE(n) sample. The median evolution from $z = 2$ to the present day prefers the minor merger–predicted evolution when comparing our study to the MASSIVE(n) sample in the dynamical–stellar mass plane.

The median evolution from our study to the MASSIVE(n) sample at the present day in the dynamical–stellar mass plane is consistent with minor merger evolution that is similar to what is found in Figures 7 and 8.

6. Discussion

The structural and kinematic evolution for massive galaxies from $z = 2$ to the present is explored by assuming that the galaxies in this study are the progenitors of the MASSIVE(n) sample. Such a claim has been motivated by a fixed CND matching between the two samples of galaxies. This suggests that these galaxies undergo significant size growth, together with shallow velocity dispersion evolution, driving up the dynamical-to-stellar mass ratio from $z = 2$ to 0. The role of major mergers in the evolution of massive galaxies is discussed following an interpretation using idealized and cosmological simulations. Furthermore, the origin of the dust heating, observed in the MIR and FIR emission, is discussed. Finally, the caveats are presented.

6.1. Quiescent-to-quiescent Major Mergers

Three galaxies in our sample, initially unresolved in ground-based imaging, are found in *HST* images to be double sources and confirmed with X-shooter to be ongoing major merger systems (see Section 4.6). In this section, we discuss how the high major merger fraction (6/18) affects the prevalence of minor merger structural evolution of MQGs (found in Section 5) and if the high fraction could be caused by a selection bias.

Following the definition in Man et al. (2012), we find a pair fraction of $20\% \pm 12\%$ when assuming no projected sources ($\langle N_{\text{projected}} \rangle = 0$) and using the Poisson error estimate. In COSMOS and UDS, a pair fraction of 10% major mergers is found for massive ($\log(M_*/M_\odot) > 11$) galaxies at $z = 2$ (Mundy et al. 2017). In the case that the observed major mergers are representative of the complete sample of MQGs, we can estimate the number of major mergers each galaxy undergoes (N_{merger}) following the prescription in Man et al. (2016b). Under the assumption that the merger rate is constant from $z = 2$ to 0, Equation (3) in Man et al. (2016b) can be written as $N_{\text{merger}} = \Delta t f_{\text{pair}} / t_{\text{obs}}$. The pair fraction $f_{\text{pair}} = 0.2$, observation time $t_{\text{obs}} = 0.8$ Gyr (from $z = 1.9$ to 2.5), and time of evolution $\Delta t \sim 10$ Gyr is used to estimate the number of mergers from $z = 2$ to 0. These numbers reproduce a major merger rate of ~ 1 for a pair fraction of 10%, similar to what was suggested in Man et al. (2012). For a 20% pair fraction, we find that, on average, each galaxy undergoes $N_{\text{merger}} = 2.5$ mergers in the range $0 < z < 2$. This number of 1:1 major mergers corresponds to a stellar mass increase of 0.5 dex, which is too high compared to the median stellar mass of the MASSIVE(n) sample (see Figure 8). In the case of a 10% pair fraction, the stellar mass increase is consistent with the average stellar mass of the MASSIVE(n) sample; however, in this case, another mechanism must then be in place to produce the large

size growth observed between our sample and the MASSIVE(n) galaxies.

Our sample was selected to be *UVJ* quiescent and *K*-band bright, which could have introduced a bias for ground-based unresolved bright red systems like quiescent-to-quiescent galaxy major mergers (see also Section 6.5). See also Mowla et al. (2018) and Marsan et al. (2019), who addressed the issue of close pairs of MQGs at $z \sim 2$. If this selection bias is responsible for the high pair fractions, this could explain why we observe that the stellar mass–size evolution from $z = 2$ to 0 is dominated by minor mergers (see Figure 8). The majority of our major merger targets are in the low stellar mass end of our sample ($\log(M_*/M_\odot) < 11.5$). This could indicate that a possible way to produce ultramassive ($\log(M_*/M_\odot) > 11.5$) QGs could be via quiescent-to-quiescent galaxy major mergers at $z > 2$. A scenario involving early-time major and late-time minor merger evolution will be testable with larger samples of MQGs at $z > 2.5$.

6.2. Minor Merger Size Evolution at Constant Dispersion

In Figure 8, a slope of $\alpha = 1.78_{-0.29}^{+0.37}$ is found for the mass–size evolution of our MQGs from $z = 2$ to 0. Such an evolution can be interpreted using the analytical framework from Bezanson et al. (2009) and Naab et al. (2009), who found that minor merger–driven growth is needed to produce a mass–size slope of $\alpha = 2$. An extended numerical treatment from Hilz et al. (2012) finds that when including the effect of escaping particles (a process arising from virialization following merger interaction), they recovered a steeper mass–size slope ($\alpha \approx 2.4$) alongside a constant dispersion evolution for minor merger–driven growth. Such a scenario could explain the observed size growth and shallow dispersion evolution observed.

The scenario presented in Hilz et al. (2012) occurs for two-component (stellar+halo) systems when they undergo 1:10 minor merger evolution. They reproduce the structural evolution found in Bezanson et al. (2009) and Naab et al. (2009) when simulating minor merger evolution of stellar-only systems. According to Hilz et al. (2012), this suggests that the growth of the dark matter halo is an important ingredient necessary to cause the shallow dispersion evolution together with the expected size growth evolution we find in this study. Moreover, Hilz et al. (2012) showed that major mergers increase the dispersion and size proportional to the stellar mass. This is not what is found when comparing the size and dispersion evolution with the MASSIVE(n) sample (see Figures 7 and 8). In the minor merger scenario, the velocity dispersion would be maintained in the inner region of the galaxy, as additional stellar mass is accreted in the outer parts from tidally stripped satellite systems. Over time, this would change the stellar light distribution on the outskirts of the galaxy, causing continuous growth of the half-light radius (van Dokkum et al. 2010; Hill et al. 2017).

In UV-105842, we may be observing a direct example of the minor merger–driven size increase. A small satellite system within close (spectroscopically confirmed) proximity of the central galaxy is found. Based on the flux ratio estimated from the GALFIT modeling, we estimate a stellar mass ratio of $1:12_{-3}^{+6}$ for this minor merger, consistent with the average 1:16 ratio estimated by Newman et al. (2012). To double its stellar mass (as suggested by the median ~ 0.3 dex increase derived for our sample), the galaxy would need to go through ~ 12 such

minor mergers between $z = 2$ and 0. Other minor merger stellar mass ratios of 1:5, 1:10, and 1:20, suggested by Hilz et al. (2013) and Bédorf & Portegies Zwart (2013), would correspond to 5, 10, and 20 minor mergers between $z = 2$ and 0 for a similar stellar mass increase. In Man et al. (2016b), issues related to the translation of the H -band flux ratio to a stellar mass ratio (e.g., due to M/L ratio variation in galaxies) directly affecting the above argument are discussed.

Many observational (Bluck et al. 2012; McLure et al. 2013; Fagioli et al. 2016; Matharu et al. 2019; Zahid et al. 2019) and numerical (Naab et al. 2009, 2014; Oser et al. 2012; Oogi & Habe 2013; Tapia et al. 2014; Remus et al. 2017) studies find that minor mergers could be a dominant process for the size growth of massive galaxies, but they may not be able to explain the full size evolution (Cimatti et al. 2012; Newman et al. 2012). Feedback processes have been shown to also affect the size growth (e.g., Lackner et al. 2012; Hirschmann et al. 2013). Specifically, AGN feedback is shown by modern simulations to be necessary to reproduce the observed size evolution (see Dubois et al. 2013; Choi et al. 2018).

6.3. Stellar-to-dynamical Mass Evolution

We found that the dynamical-to-stellar mass ratio shown in Figure 9 increases by a factor of 2 within MQGs from $z = 2$ to 0. This could be attributed to either IMF changes of the stellar population (Cappellari et al. 2012) affecting the stellar mass estimates or an increase in the dark matter fraction within the effective half-light radius.

Numerical simulations find that minor merger-driven evolution alters the distribution of stars over time from a core to a core-envelope system by accretion of particles in the outskirts of the galaxy (Hopkins et al. 2009; Hilz et al. 2012, 2013; Frigo & Balcels 2017; Lagos et al. 2018). A consequence of this is that the central dispersion remains constant while the half-light radius grows, encompassing a larger part of the dark matter halo and effectively increasing the dark matter fraction over time (Hilz et al. 2012).

A mass-size evolution similar to what we find is, according to Hilz et al. (2013), caused by a massive dark matter halo that drives the accretion of dry (collisionless) minor mergers at large radii through tidal stripping. This inside-out growth increases the effective half-mass radius to encompass dark matter-dominated regions, which might explain the increase of the dynamical-to-stellar mass fraction within the half-light radius that we observe.

Care must be taken when interpreting the observations in terms of idealized numerical simulations. However, Remus et al. (2017) also found that the central dark matter fraction increases with decreasing redshift when comparing different cosmological simulations. Furthermore, observational evidence for inside-out growth in massive galaxies is presented in Szomoru et al. (2012).

In Figure 9, we find that our sample is consistent with the dynamical-to-stellar mass ratio of one suggesting low dark matter fractions at $z \sim 2$. For a stellar mass increase of 0.3 dex (similar to our median evolution), Hilz et al. (2012) predicted a dark matter fraction increase of $\sim 70\%$ within the effective radius. If we assume that the mass of the galaxy consists only of dark matter and stars, we can estimate the dark matter mass fractions ($M_{\text{DM}}/M_{\text{dyn}} = 1 - M_*/M_{\text{dyn}}$) from the dynamical-to-stellar median ratio at $z = 2$ and 0 to be $7_{-7}^{+24}\%$ and $56\% \pm 8\%$, respectively. This suggests an increase of the dark matter

fraction within the effective radius of 17%–64%. Note, however, that this increase cannot purely be associated with the dark matter from the minor mergers, as the growing half-light radius similarly encompasses more of the central dark matter halo that therefore also contributes to this increase.

According to Remus et al. (2017), the mass growth of massive galaxies can be explained by two stages: (1) high-redshift in situ mass growth resulting in a dense stellar component in the center of the potential where the dark matter fraction is low and (2) dry merger events dominating the mass growth at lower redshift (with major mergers being rare) resulting in the buildup of a stellar envelope increasing the half-light radius and thus the dark matter fraction (similar to the interpretation above).

6.4. Dust Heating in MQGs at $z > 2$

The $24\ \mu\text{m}$ SFR limit, used to restrict the stellar population models, results in sSFRs for our galaxies of $\log_{10}(\text{sSFR}/\text{yr}) < -10$. Nonetheless, stronger limits on the sSFR can be obtained if the source of dust heating is not caused by recent SF. In Section 4.4.3, the information from optical nebular emission and MIR is combined to set stringent limits on the SFR of our sample (see also Figure 4). This information reveals that our sample lies 1.5 dex below the $z = 2$, SFR-stellar mass relation mass relation of Speagle et al. (2014; extrapolated to $\log_{10}(M_{\odot}/M_*) \sim 11.5$).

Low-luminosity AGNs are shown to be common in massive, $\log_{10}(M_*/M_{\odot}) > 11$, quiescent galaxies at $z < 1.5$ through excess radio emission in stacked samples (Man et al. 2016a; Gobat et al. 2018). Six galaxies in our sample have direct radio detections, three of them with matching MIR detections (see Table 2). This could be evidence in line with the results from Olsen et al. (2013), who found a high fraction of AGNs in MQGs at $1.5 < z < 2.5$ using X-ray stacking. Low-luminosity AGN activity has, in Schawinski et al. (2009) and Best & Heckman (2012), been associated with the suppression of SF, which is an important effect in maintaining quiescent galaxies. Low levels of dust heating have also been associated with evolved stellar populations as a significant source to emit at wavelengths beyond $> 160\ \mu\text{m}$ (Salim et al. 2009; Bendo et al. 2012; Fumagalli et al. 2014; Utomo et al. 2014). However, with no detections in the *Herschel*/PACS bands, we cannot rule this scenario out. In the case where AGNs are indeed the dominant dust heating source in the galaxies, we can expect that the $24\ \mu\text{m}$ flux does not arise from residual SF. This is consistent with Whitaker et al. (2017), who found no strongly obscured SF in MQGs at $z > 2$. Assuming the $24\ \mu\text{m}$ emission is not due to obscured SF, we find an sSFR, $\log_{10}(\text{sSFR}/\text{yr}) < -11$, based purely on the optical emission limits/detections. The MIR-to-radio emission of the sample will, in a future publication, be investigated in detail (I. Cortzen et al. 2020, in preparation).

6.5. Caveats

The main limitations of the results are presented as follows.

1. Overestimated stellar masses would lead to a shallower mass-size evolution and dynamical-to-stellar mass ratio evolution. Nonetheless, substantially overestimated stellar masses are ruled out by our dynamical masses being in agreement with previous kinematic studies of MQGs at $z > 2$ (Toft et al. 2012; Bezanson et al. 2013; van de Sande et al. 2013; Belli et al. 2014b, 2017).

2. If rotation is significant in MQGs at $z > 2$, the measured velocity dispersion, depending on the inclination, could have an unknown contribution from rotation resulting in heightened dispersion measurements. On the other hand, dispersion measurements from face-on rotation-dominated galaxies could result in low values. This would further drive the dynamical mass artificially down. Such issues should be addressed by spatially resolved spectroscopy where the V_{rot}/σ can be estimated.
3. Previous studies (Mancini et al. 2010) have suggested that sizes might be underestimated due to nondetection of low-luminosity profile wings. However, ultradeep imaging out to many effective radii does not find that this is the case (Szomoru et al. 2010, 2011).
4. Dynamical-to-stellar mass evolution is sensitive to the determination of $\beta(n)$. The prescription from Cappellari et al. (2006) is used, yet this relation is determined from local galaxies and assumed to be representative of dynamical systems at $z \sim 2$. When comparing with the MASSIVE(n) sample, we assume a Sérsic index of $n = 4$ to be a fair representation of a spheroidal system. When changing the choice of $\beta = 2-6$ for the MASSIVE(n) sample, the conclusion that the ratio must evolve from $z = 2$ to 0 remains.
5. The sample is 60% mass complete for the massive ($\log_{10}(M_*/M_\odot) > 11$) and K -band-brightest ($K < 20.5$) UVJ quiescent galaxies at $1.9 < z < 2.5$. This selection depends strongly on the performance of the photometric redshift estimate. In Figure 3, we show that this works well for our sample using the catalog from Muzzin et al. (2013a). This suggests that the sample studied in this paper is representative of the selection we presented in Section 2. However, the photometry is used to select red systems and, consequently, introduce a selection bias toward mergers between red galaxies. An unresolved merger of a quiescent galaxy with a star-forming galaxy would produce a resulting bluer system that might be excluded from the selection (see Section 6.1).

7. Summary and Conclusion

We examined the largest sample of MQGs observed to date at $z > 2$ with deep X-shooter spectroscopy and *HST*/WFC3 imaging. We extend previous searches for very MQGs at $z > 2$ to the K -band-brightest UVJ quiescent galaxies in COSMOS (Muzzin et al. 2013a), constructing a sample of 15 MQGs. Full SED modeling of the photometry and spectroscopy confirms the sample to be ~ 1.5 Gyr old, massive, $\log_{10}(M_*/M_\odot) > 11$, quiescent galaxies. Three out of 15 galaxies are confirmed as ongoing major mergers using both imaging and spectroscopy. In total, 40% of the sample shows evidence of mergers (minor or major) or other disturbed morphologies in *HST*/WFC3 H_{F160W} imaging, suggestive of ongoing morphological transformation. The morphological information is used to correct the stellar masses prior to comparing the stellar populations, kinematics, and structure/morphology of the galaxies to the MASSIVE(n) sample. Below, we list the main conclusions of the paper.

1. We find that our galaxies lie 1–1.5 dex below the extrapolation at the high stellar mass end of the SFR main sequence (Speagle et al. 2014) at $z = 2$ and can be

considered quiescent with low sSFR, $\log_{10}(\text{sSFR}/\text{yr}) < -10.5$. These limits are based on optical emission line and MIR emission limits and detections. One-third of the galaxies are detected in the MIR, which could be caused by residual SF. However, more than half of our sample (60% of the MIR detections) has radio emission detected at 1.4 or 3 GHz. This radio emission is likely associated with AGN activity, a proposed heating mechanism leading to quenching and/or the maintenance of quiescence in massive galaxies.

2. We find indirect evidence pointing to our velocity dispersion measurements being minimally contaminated by rotation. Our systems also have a Sérsic index $n > 2.5$ (see Section 5.2). A direct comparison between our study and the MASSIVE(n) sample shows evidence for shallow or no velocity dispersion evolution from $z = 2$ to 0.
3. Our sample is compact, in line with previous studies at $z \sim 2$ (van der Wel et al. 2014; Mowla et al. 2018). We find that the median mass–size evolution ($\Delta r \propto \Delta M_*^\alpha$) compared to the MASSIVE(n) sample is best described by $\alpha = 1.78^{+0.37}_{-0.29}$. This is consistent with both the simple kinematic predictions of minor merger–driven size evolution from Bezanson et al. (2009) and the more extensive numerical treatment from Hilz et al. (2012).
4. We find that our sample of $z > 2$ MQGs is consistent with a dynamical-to-stellar mass ratio $M_*/M_{\text{dyn}} < 1$ but that the shallow dispersion and significant size increase lead to an increasing dynamical-to-stellar mass ratio, doubling from $z = 2$ to the present day. Such an effect is shown to be reproduced for an increasing dark matter fraction from $z = 2$ to 0, within the effective radius of the galaxy (Hilz et al. 2012).

In this paper, the largest sample of MQGs at $z > 2$ with kinematic and structural observations, found via the mass–size and dynamical–stellar mass plane, is presented. A fixed CND matching suggests that the galaxies in our sample are the progenitors of the most massive and oldest elliptical galaxies in the local universe, thus connecting 10 billion yr of evolution. These galaxies show a broad range of disturbed morphologies, confirming that mergers play a significant role in their morphological transformation and evolution to $z = 0$.

In a companion paper, the relationship between the size and dispersion will be explored by studying the fundamental plane at $z \sim 2$ and its consequent evolution to the present-day universe (M. Stockmann et al. 2020, in preparation).

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Appendix A

Further Details on the Reduction of the Images

A.1. PSF and Astrometry

The H_{F160W} images from our program and the ancillary COSMOS F814W images employed in this work do not share the same world coordinate system. We need to guarantee that the astrometry is common and accurate in both bands. Therefore, we choose to align the images to the COSMOS ACS F814W image as the reference frame, which is registered to the fundamental astrometric frame of the COSMOS field, ensuring an absolute astrometric accuracy of $0''.05$ – $0''.1$ or better. Following Gómez-Guijarro et al. (2018), we use *TweakReg* along with *SExtractor* (Bertin & Arnouts 1996) catalogs of the two bands with the F814W catalog and frame as references to register the images. After this, the images in both bands are resampled to a common grid and a pixel scale of $0''.06 \text{ pixel}^{-1}$ using *SWarp* (Bertin et al. 2002). In addition, the spatial resolution of the two *HST* bands is also different. Following Gómez-Guijarro et al. (2018), we degrade the F814W to the resolution of the F160W data ($0''.18$ FWHM). We calculate the kernel to match the ACS F814W to the PSF in the F160W images employing the task *PSFMATCH* in *IRAF*, including a cosine bell function tapered in frequency space to avoid introducing artifacts in the resulting kernel from the highest frequencies. Then, we convolve this kernel to the F814W image to achieve a common spatial resolution.

A.2. Modeling of Foreground and Background Sources

Based on the spatially offset emission in the 2D X-shooter spectra, we determine whether candidate sources are within close proximity to the central galaxy. In Figure 5, we show the

galaxies together with their spatially offset companions. The object UV-171687 shows offset $H\alpha$ and $[N II]$ emission arising from a southwestern source that we establish to be a foreground galaxy at $z = 1.51$. We find another foreground galaxy northeast of UV-171060 at $z = 1.37$ based on assuming that the single emission line detection is $H\alpha$. Northeast of UV-155853, we find a background galaxy at $z = 2.36$ (best visible in the GALFIT modeling residuals of Figure 5) determined from the $[O III]$ doublet at $4959,5007 \text{ \AA}$. For UV-105842, we find two spatially offset sources (1) $\sim 3''$ northeast and (2) $\sim 1''$ northeast. Source (1) is a foreground galaxy at $z = 0.44$ based on the detected strong $[O III]$ doublet at $4959,5007 \text{ \AA}$ and $H\alpha$ emission. For source (2), we find the $[O II]$ doublet at $3726.2, 3728.9 \text{ \AA}$, $[O III]$ doublet at $4959,5007 \text{ \AA}$, and $H\alpha$ corresponding to a redshift $z = 2.0124$. The latter redshift corresponds to a velocity offset of $2130 \pm 120 \text{ km s}^{-1}$ (uncertainty is calculated based on the spread of the individual redshift measurements), suggesting that it is not gravitationally bound to the central galaxy. Another option could be an offset AGN with high peculiar velocity following a merger.

Appendix B

Details on Modeling of the Velocity Dispersion

B.1. Statistical and Systematic Uncertainties

To estimate the statistical error, we measure the spread of the velocity dispersion distribution obtained from running pPXF on 1000 data realizations. The data realizations are made by perturbing the pPXF best-fit model with the pipeline-estimated error spectrum by linearly drawing values from a Gaussian with a mean of zero and spread of the initial errors. The X-shooter pipeline-estimated noise map is subjected to a wavelength-dependent correlation of the pixels. We take this effect into account by scaling our noise spectrum to a reduced $\chi^2_{\text{red}} = 1$ (assuming the errors are Gaussian). We follow the method used in Toft et al. (2017) and fit a second-order polynomial to a 50 pixel running reduced χ^2 that we use to make a correction noise map, $\sigma_{\chi^2_{\text{corr}}} = \sigma_{\chi^2_{\text{original}}} \sqrt{\chi^2_{\text{fit}}}$.

We estimate the systematic error by testing how the dispersion is changing with the correction polynomial and implemented wavelength range. We construct a grid of correction polynomials up to 24th order of both additive and multiplicative polynomials, where we find an average of 20% variation from the fiducial dispersion, except for UV-239220, UV-773654, UV-171060, and CP-1291751. When varying the starting wavelength range ($[\lambda_{\text{start}}, \lambda_{\text{end}}]$) within the interval $[3750\text{--}4050, 5950]$ and the ending wavelength within the interval $[3750, 4050\text{--}5950]$, we find that overall, the dispersions are stable. In a few cases, the velocity dispersion increases well above the median dispersion (with varying wavelengths) with 50%–100% when excluding the higher-order Balmer and Ca H+K lines, highlighting their importance. When including the ending wavelength $\lambda > 4500$, we find more stable dispersion measurements; this is not surprising, as otherwise only half of the spectrum is included. The low-S/N cases have more unstable dispersion values when excluding wavelength areas, highlighting the importance of understanding the systematic uncertainties. We sum up the wavelength and polynomial test by confirming that our fiducial velocity dispersions are robust (except for UV-239220, UV-773654, UV-171060, and CP-1291751). The systematic error is primarily due to template mismatch, and as a result, we

²⁸ <http://www.astropy.org>

estimate the systematic error from the minimum and maximum values of the dispersion when using the full wavelength range and varying the additive and multiplicative correction polynomials, $\sigma_{\text{sys}} = 2/3 \cdot (\sigma_{\text{max}} - \sigma_{\text{min}})/2$. This method is subjected to catastrophic outliers, and prior to the systematic error estimate, we exclude dispersion values more than 5σ outside of a Gaussian mean. We find that the systematic errors are on the order of the statistical uncertainties.

B.1.1. Additional Tests

We measure the dispersion while excluding a window of 1600 km s^{-1} along the wavelength direction in steps of 5 \AA to test whether the measured dispersion is dominated by specific lines. We find that the fiducial dispersion is very stable against excluding individual lines and do not find a consistent decrease in the velocity dispersion similar to previous studies when excluding the $\text{H}\beta$ line (van de Sande et al. 2013; Toft et al. 2017). We allow pPXF to construct a linear combination of templates from the stellar library of BC03 with a Chabrier IMF and solar metallicity and find similar redshifts and velocity dispersions as our fiducial values, which is reassuring.

Appendix C Details on the Emission Line Fitting

For UV-108899, we find that fitting a double Gaussian profile to $[\text{O II}]$ ($3726 + 3729 \text{ \AA}$) fixed to the redshift of the central galaxy gives the most conservative (highest) flux estimate. We try fitting with a single profile while using a free redshift parameter but recover high χ^2 solutions. We list this conservative flux estimate, corresponding to an $\text{SFR} = 6 \pm 4 M_{\odot} \text{ yr}^{-1}$ (Kennicutt 1998), in Table 2.

For UV-239220, we detect excess emission in the region of $\text{H}\alpha$ and the $[\text{N II}]$ ($6548 + 6583 \text{ \AA}$) doublet. With a fixed ratio between the $[\text{N II}]$ doublet, we try three types of triple Gaussian profile models (free redshift+dispersion limit of 250 km s^{-1} , free redshift+dispersion limit of 1000 km s^{-1} , and fixed redshift+dispersion limit of 1000 km s^{-1}) that all result in $\chi^2 > 2.4$ with no preferred solution. If we assign all of the flux in the excess to $\text{H}\alpha$, we obtain a conservative Kennicutt (1998) SFR upper limit of $\sim 30 M_{\odot} \text{ yr}^{-1}$ ($\log(\text{sSFR}) < -10 [\text{yr}^{-1}]$), consistent with the FIR and rest-frame optical upper limits from Section 4.4.1. This confirms that the galaxy has a low specific SFR, consistent with its selection.

Appendix D Comparing Different CND Methods

The MASSIVE(n) sample was established using the assumption of a fixed CND from $z = 2$ to 0. To show that our results are robust against the choice of CND-matching method, we show the three result figures from Section 5 in Figure 10 using both the fixed CND matching and the probabilistic approach presented in Wellons & Torrey (2017).

The probabilistic approach uses numerical simulations (e.g., Illustris) to estimate the probability that a galaxy at $z = 0$ is the descendant of a galaxy at redshift, z_{obs} . This method therefore allows one to predict the most probable CND at $z = 0$ for a population of galaxies with the specific CND at $z = 2$ following the evolution of a numerical simulation. This method is thus a different approach than the fixed CND approach, and in Figure 10, we show that adopting these two methods of connecting galaxies across time leads to the same conclusions.

Fixed CDN

Probabilistic CDN

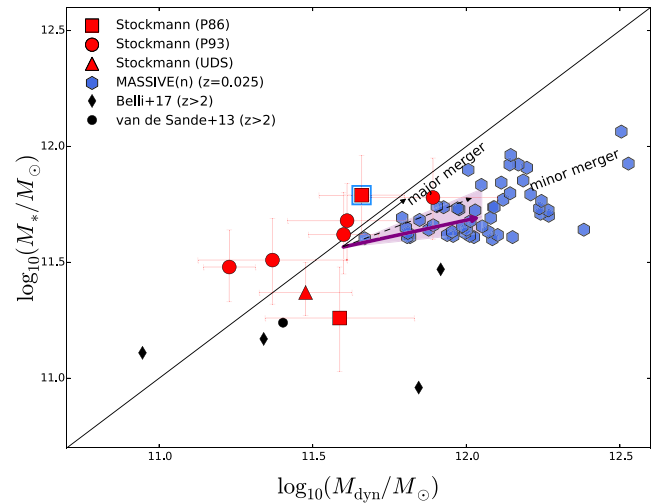
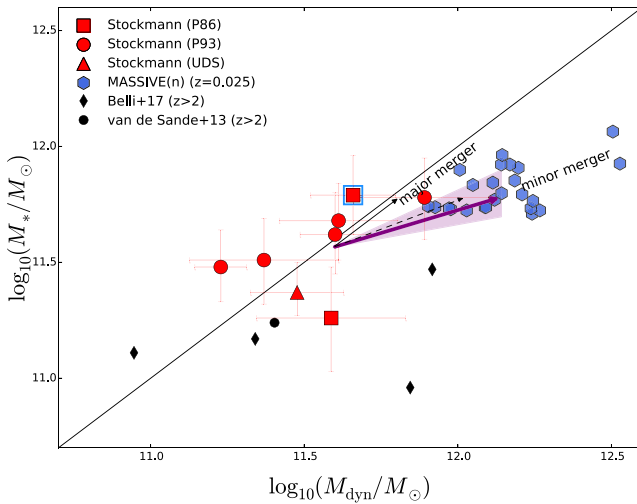
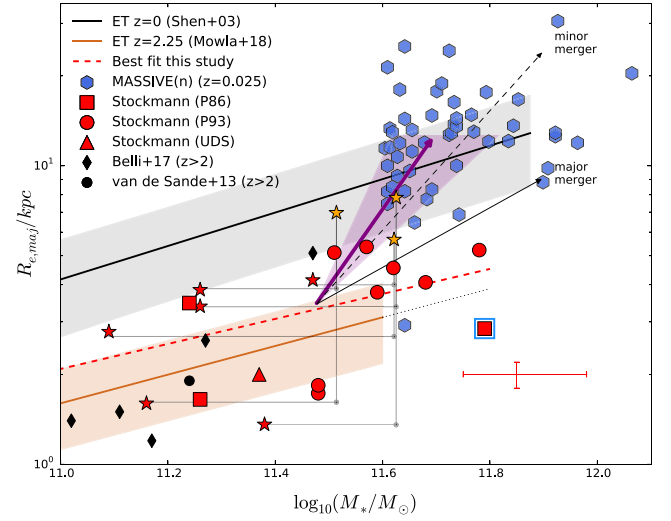
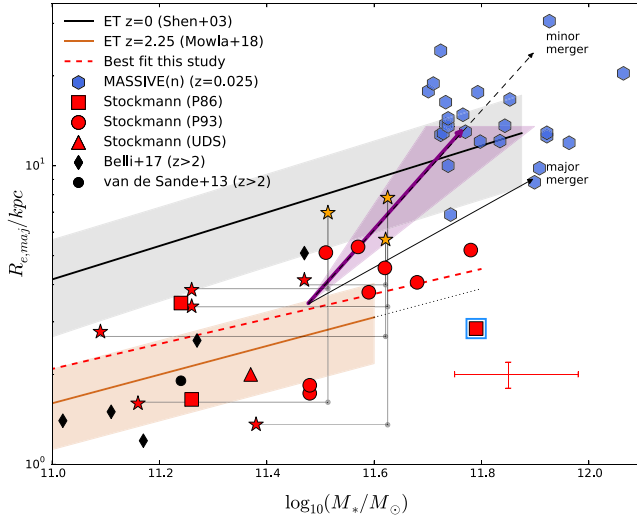
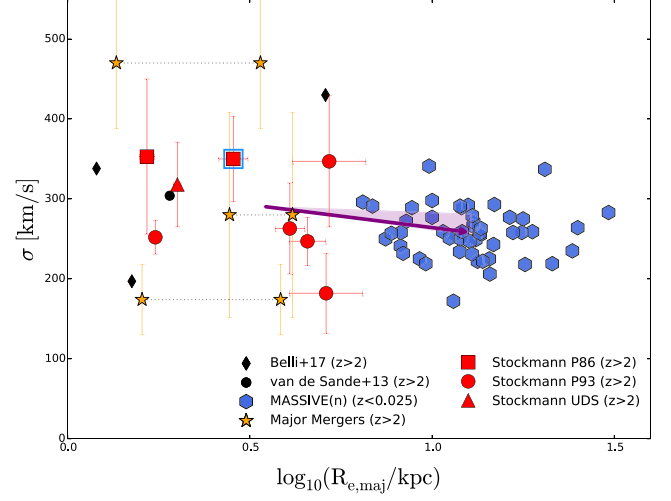
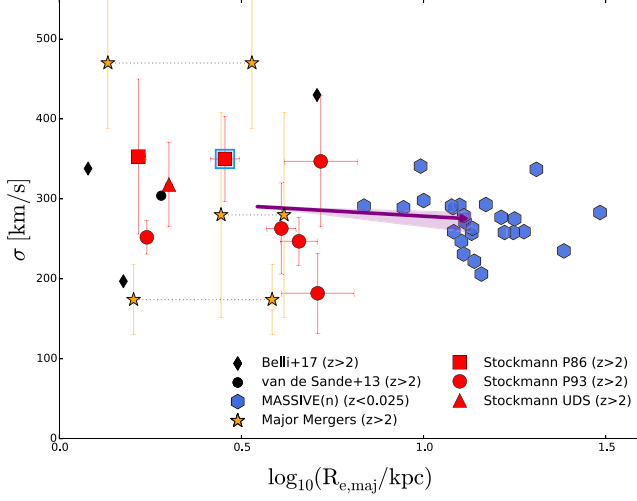


Figure 10. The figures from Section 5 are shown with the fixed and probabilistic (Wellons & Torrey 2017) CDN match to the MASSIVE Survey. For each of these methods, our qualitative conclusions remain.

ORCID iDs

Mikkel Stockmann  <https://orcid.org/0000-0001-5983-6273>
 Sune Toft  <https://orcid.org/0000-0003-3631-7176>
 Anna Gallazzi  <https://orcid.org/0000-0002-9656-1800>
 Stefano Zibetti  <https://orcid.org/0000-0003-1734-8356>
 Christopher J. Conselice  <https://orcid.org/0000-0003-1949-7638>
 Berta Margalef-Bentabol  <https://orcid.org/0000-0001-8702-7019>
 Johannes Zabl  <https://orcid.org/0000-0002-9842-6354>
 Inger Jørgensen  <https://orcid.org/0000-0003-3002-1446>
 Georgios E. Magdis  <https://orcid.org/0000-0002-4872-2294>
 Carlos Gómez-Guijarro  <https://orcid.org/0000-0002-4085-9165>
 Francesco M. Valentino  <https://orcid.org/0000-0001-6477-4011>
 Gabriel B. Brammer  <https://orcid.org/0000-0003-2680-005X>
 Daniel Ceverino  <https://orcid.org/0000-0002-8680-248X>
 Isabella Cortzen  <https://orcid.org/0000-0001-9197-7623>
 Iary Davidzon  <https://orcid.org/0000-0002-2951-7519>
 Richardo Demarco  <https://orcid.org/0000-0003-3921-2177>
 Andreas Faisst  <https://orcid.org/0000-0002-9382-9832>
 Michaela Hirschmann  <https://orcid.org/0000-0002-3301-3321>
 Jens-Kristian Krogager  <https://orcid.org/0000-0002-4912-9388>
 Claudia D. Lagos  <https://orcid.org/0000-0003-3021-8564>
 Allison W. S. Man  <https://orcid.org/0000-0003-2475-124X>
 Carl J. Mundy  <https://orcid.org/0000-0002-0833-8554>
 Yingjie Peng  <https://orcid.org/0000-0001-8302-5198>
 Jonatan Selsing  <https://orcid.org/0000-0001-9058-3892>
 Charles L. Steinhardt  <https://orcid.org/0000-0003-3780-6801>
 Kathrine E. Whitaker  <https://orcid.org/0000-0001-7160-3632>

References

- Allen, R. J., Kacprzak, G. G., Spitler, L. R., et al. 2015, *ApJ*, 806, 3
- Arcila-Osejo, L., Sawicki, M., Arnouts, S., et al. 2019, *MNRAS*, 486, 4880
- Astropy Collaboration, Price-Whelan, A. M., Sipőcz, B. M., et al. 2018, *AJ*, 156, 123
- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, 558, A33
- Barden, M., Häußler, B., Peng, C. Y., McIntosh, D. H., & Guo, Y. 2012, *MNRAS*, 422, 449
- Bédorf, J., & Portegies Zwart, S. 2013, *MNRAS*, 431, 767
- Behroozi, P. S., Marchesini, D., Wechsler, R. H., et al. 2013, *ApJL*, 777, L10
- Bell, E. F., van der Wel, A., Papovich, C., et al. 2012, *ApJ*, 753, 167
- Belli, S., Newman, A. B., & Ellis, R. S. 2014a, *ApJ*, 783, 117
- Belli, S., Newman, A. B., & Ellis, R. S. 2017, *ApJ*, 834, 18
- Belli, S., Newman, A. B., & Ellis, R. S. 2018, *ApJ*, 874, 17
- Belli, S., Newman, A. B., Ellis, R. S., & Konidaris, N. P. 2014b, *ApJL*, 788, L29
- Bendo, G. J., Boselli, A., Dariush, A., et al. 2012, *MNRAS*, 419, 1833
- Bertin, E., & Arnouts, S. 1996, *A&AS*, 117, 393
- Bertin, E., Mellier, Y., Radovich, M., et al. 2002, in ASP Conf. Ser. 281, Astronomical Data Analysis Software and Systems XI, ed. D. A. Bohlender, D. Durand, & T. H. Handley (San Francisco, CA: ASP), 228
- Best, P. N., & Heckman, T. M. 2012, *MNRAS*, 421, 1569
- Bezanson, R., van Dokkum, P., & Franx, M. 2012, *ApJ*, 760, 62
- Bezanson, R., van Dokkum, P. G., Tal, T., et al. 2009, *ApJ*, 697, 1290
- Bezanson, R., van Dokkum, P. G., van de Sande, J., et al. 2013, *ApJL*, 779, L21
- Bezanson, R., van der Wel, A., Pacifici, C., et al. 2018, *ApJ*, 858, 60
- Bluck, A. F. L., Conselice, C. J., Buitrago, F., et al. 2012, *ApJ*, 747, 34
- Brammer, G. B., van Dokkum, P. G., & Coppi, P. 2008, *ApJ*, 686, 1503
- Brammer, G. B., Whitaker, K. E., van Dokkum, P. G., et al. 2011, *ApJ*, 739, 24
- Bruzual, G., & Charlot, S. 2003, *MNRAS*, 344, 1000
- Buitrago, F., Trujillo, I., Conselice, C. J., et al. 2008, *ApJL*, 687, L61
- Buitrago, F., Trujillo, I., Conselice, C. J., & Häußler, B. 2013, *MNRAS*, 428, 1460
- Capak, P., Aussel, H., Ajiki, M., et al. 2007, *ApJS*, 172, 99
- Cappellari, M., Bacon, R., Bureau, M., et al. 2006, *MNRAS*, 366, 1126
- Cappellari, M., & Emsellem, E. 2004, *PASP*, 116, 138
- Cappellari, M., McDermid, R. M., Alatalo, K., et al. 2012, *Natur*, 484, 485
- Cappellari, M., Scott, N., Alatalo, K., et al. 2013, *MNRAS*, 432, 1709
- Carollo, C. M., Bschorr, T. J., Renzini, A., et al. 2013, *ApJ*, 773, 112
- Chabrier, G. 2003, *PASP*, 115, 763
- Charlot, S., & Fall, S. M. 2000, *ApJ*, 539, 718
- Choi, E., Somerville, R. S., Ostriker, J. P., Naab, T., & Hirschmann, M. 2018, *ApJ*, 866, 91
- Cimatti, A., Cassata, P., Pozzetti, L., et al. 2008, *A&A*, 482, 21
- Cimatti, A., Daddi, E., Renzini, A., et al. 2004, *Natur*, 430, 184
- Cimatti, A., Nipoti, C., & Cassata, P. 2012, *MNRAS*, 422, L62
- Comerford, J. M., & Greene, J. E. 2014, *ApJ*, 789, 112
- Conselice, C. J., Blackburne, J. A., & Papovich, C. 2005, *ApJ*, 620, 564
- Conselice, C. J., Bluck, A. F. L., Buitrago, F., et al. 2011, *MNRAS*, 413, 80
- Daddi, E., Cimatti, A., Renzini, A., et al. 2004, *ApJL*, 600, L127
- Daddi, E., Renzini, A., Pirzkal, N., et al. 2005, *ApJ*, 626, 680
- Damjanov, I., Abraham, R. G., Glazebrook, K., et al. 2011, *ApJL*, 739, L44
- Djorgovski, S., & Davis, M. 1987, *ApJ*, 313, 59
- D'Odorico, S., Dekker, H., Mazzoleni, R., et al. 2006, *Proc. SPIE*, 6269, 626933
- Dressler, A., Lynden-Bell, D., Burstein, D., et al. 1987, *ApJ*, 313, 42
- Dubois, Y., Gavazzi, R., Peirani, S., & Silk, J. 2013, *MNRAS*, 433, 3297
- Fagioli, M., Carollo, C. M., Renzini, A., et al. 2016, *ApJ*, 831, 173
- Faisst, A. L., Carollo, C. M., Capak, P. L., et al. 2017, *ApJ*, 839, 71
- Falcón-Barroso, J., Sánchez-Blázquez, P., Vazdekis, A., et al. 2011, *A&A*, 532, A95
- Ferré-Mateu, A., Vazdekis, A., Trujillo, I., et al. 2012, *MNRAS*, 423, 632
- Franx, M., Labbé, I., Rudnick, G., et al. 2003, *ApJL*, 587, L79
- Franx, M., van Dokkum, P. G., Förster Schreiber, N. M., et al. 2008, *ApJ*, 688, 770
- Frigo, M., & Balcells, M. 2017, *MNRAS*, 469, 2184
- Fukugita, M., Ichikawa, T., Gunn, J. E., et al. 1996, *AJ*, 111, 1748
- Fumagalli, M., Labbé, I., Patel, S. G., et al. 2014, *ApJ*, 796, 35
- Gallazzi, A., Charlot, S., Brinchmann, J., White, S. D. M., & Tremonti, C. A. 2005, *MNRAS*, 362, 41
- Geier, S., Richard, J., Man, A. W. S., et al. 2013, *ApJ*, 777, 87
- Genel, S., Vogelsberger, M., Springel, V., et al. 2014, *MNRAS*, 445, 175
- Gobat, R., Daddi, E., Magdis, G., et al. 2018, *NatAs*, 2, 239
- Gómez-Guijarro, C., Toft, S., Karim, A., et al. 2018, *ApJ*, 856, 121
- Graves, G. J., Faber, S. M., & Schiavon, R. P. 2009, *ApJ*, 698, 1590
- Greene, J. E., Janish, R., Ma, C.-P., et al. 2015, *ApJ*, 807, 11
- Hill, A. R., Muzzin, A., Franx, M., et al. 2017, *ApJ*, 837, 147
- Hilz, M., Naab, T., Ostriker, J. P., et al. 2012, *MNRAS*, 425, 3119
- Hilz, M., Naab, T., & Ostriker, J. P. 2013, *MNRAS*, 429, 2924
- Hirschmann, M., Naab, T., Davé, R., et al. 2013, *MNRAS*, 436, 2929
- Hopkins, P. F., Bundy, K., Murray, N., et al. 2009, *MNRAS*, 398, 898
- Hopkins, P. F., Richards, G. T., & Hernquist, L. 2007, *ApJ*, 654, 731
- Horne, K. 1986, *PASP*, 98, 609
- Husser, T.-O., Wende-von Berg, S., Dreizler, S., et al. 2013, *A&A*, 553, A6
- Ilbert, O., Capak, P., Salvato, M., et al. 2009, *ApJ*, 690, 1236
- Jeans, J. H. 1902, *RSPTA*, 199, 1
- Jin, S., Daddi, E., Liu, D., et al. 2018, *ApJ*, 864, 56
- Kado-Fong, E., Marchesini, D., Marsan, Z. C., et al. 2017, *ApJ*, 838, 57
- Kennicutt, R. C., Jr. 1998, *ARA&A*, 36, 189
- Kennicutt, R. C., Jr., Hao, C.-N., Calzetti, D., et al. 2009, *ApJ*, 703, 1672
- Khochfar, S., & Silk, J. 2006, *ApJL*, 648, L21
- Koekemoer, A. M., Aussel, H., Calzetti, D., et al. 2007, *ApJS*, 172, 196
- Kriek, M., Conroy, C., van Dokkum, P. G., et al. 2016, *Natur*, 540, 248
- Kriek, M., van Dokkum, P. G., Franx, M., et al. 2006a, *ApJ*, 645, 44
- Kriek, M., van Dokkum, P. G., Franx, M., et al. 2006b, *ApJL*, 649, L71
- Kriek, M., van Dokkum, P. G., Labbé, I., et al. 2009, *ApJ*, 700, 221
- Krogager, J.-K., Zirm, A. W., Toft, S., Man, A., & Brammer, G. 2014, *ApJ*, 797, 17
- Labbé, I., Huang, J., Franx, M., et al. 2005, *ApJL*, 624, L81
- Lackner, C. N., Cen, R., Ostriker, J. P., & Joung, M. R. 2012, *MNRAS*, 425, 641
- Lagos, C. d. P., Stevens, A. R. H., Bower, R. G., et al. 2018, *MNRAS*, 473, 4956
- Laigle, C., McCracken, H. J., Ilbert, O., et al. 2016, *ApJS*, 224, 24
- Le Floc'h, E., Aussel, H., Ilbert, O., et al. 2009, *ApJ*, 703, 222
- Leja, J., van Dokkum, P., & Franx, M. 2013, *ApJ*, 766, 33
- Ma, C.-P., Greene, J. E., McConnell, N., et al. 2014, *ApJ*, 795, 158
- Madau, P., & Dickinson, M. 2014, *ARA&A*, 52, 415
- Magdis, G. E., Daddi, E., Béthermin, M., et al. 2012, *ApJ*, 760, 6
- Man, A. W. S., Greve, T. R., Toft, S., et al. 2016a, *ApJ*, 820, 11

- Man, A. W. S., Toft, S., Zirm, A. W., Wuyts, S., & van der Wel, A. 2012, *ApJ*, **744**, 85
- Man, A. W. S., Zirm, A. W., & Toft, S. 2016b, *ApJ*, **830**, 89
- Mancini, C., Daddi, E., Renzini, A., et al. 2010, *MNRAS*, **401**, 933
- Marchesi, S., Civano, F., Elvis, M., et al. 2016, *ApJ*, **817**, 34
- Marchesini, D., Muzzin, A., Stefanon, M., et al. 2014, *ApJ*, **794**, 65
- Marsan, Z. C., Marchesini, D., Muzzin, A., et al. 2019, *ApJ*, **871**, 201
- Matharu, J., Muzzin, A., Brammer, G. B., et al. 2019, *MNRAS*, **484**, 595
- McCracken, H. J., Capak, P., Salvato, M., et al. 2010, *ApJ*, **708**, 202
- McLure, R. J., Pearce, H. J., Dunlop, J. S., et al. 2013, *MNRAS*, **428**, 1088
- Modigliani, A., Goldoni, P., Royer, F., et al. 2010, *Proc. SPIE*, **7737**, 773728
- Morishita, T., Abramson, L. E., Treu, T., et al. 2018, *ApJ*, **877**, 141
- Mowla, L., van Dokkum, P., Brammer, G., et al. 2018, *ApJ*, **880**, 57
- Mundy, C. J., Conselice, C. J., Duncan, K. J., et al. 2017, *MNRAS*, **470**, 3507
- Mundy, C. J., Conselice, C. J., & Ownsworth, J. R. 2015, *MNRAS*, **450**, 3696
- Muzzin, A., Marchesini, D., Stefanon, M., et al. 2013a, *ApJS*, **206**, 8
- Muzzin, A., Marchesini, D., Stefanon, M., et al. 2013b, *ApJ*, **777**, 18
- Naab, T., Johansson, P. H., & Ostriker, J. P. 2009, *ApJL*, **699**, L178
- Naab, T., Khochfar, S., & Burkert, A. 2006, *ApJL*, **636**, L81
- Naab, T., Oser, L., Emsellem, E., et al. 2014, *MNRAS*, **444**, 3357
- Nelson, D., Pillepich, A., Genel, S., et al. 2015, *A&C*, **13**, 12
- Newman, A. B., Belli, S., Ellis, R. S., & Patel, S. G. 2018, *ApJ*, **862**, 126
- Newman, A. B., Ellis, R. S., Bundy, K., & Treu, T. 2012, *ApJ*, **746**, 162
- Oke, J. B., & Gunn, J. E. 1983, *ApJ*, **266**, 713
- Olsen, K. P., Rasmussen, J., Toft, S., & Zirm, A. W. 2013, *ApJ*, **764**, 4
- Oogi, T., & Habe, A. 2013, *MNRAS*, **428**, 641
- Oser, L., Naab, T., Ostriker, J. P., & Johansson, P. H. 2012, *ApJ*, **744**, 63
- Papovich, C., Dickinson, M., Giallisco, M., Conselice, C. J., & Ferguson, H. C. 2005, *ApJ*, **631**, 101
- Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2002, *AJ*, **124**, 266
- Remus, R.-S., Dolag, K., Naab, T., et al. 2017, *MNRAS*, **464**, 3742
- Rieke, G. H., Alonso-Herrero, A., Weiner, B. J., et al. 2009, *ApJ*, **692**, 556
- Robitaille, T., & Bressert, E. 2012, APLpy: Astronomical Plotting Library in Python, v 0.9.12, Astrophysics Source Code Library, ascl:1208.017
- Salim, S., Dickinson, M., Michael Rich, R., et al. 2009, *ApJ*, **700**, 161
- Sánchez-Blázquez, P., Peletier, R. F., Jiménez-Vicente, J., et al. 2006, *MNRAS*, **371**, 703
- Sandage, A. 1986, *A&A*, **161**, 89
- Schawinski, K., Lintott, C. J., Thomas, D., et al. 2009, *ApJ*, **690**, 1672
- Scoville, N., Aussel, H., Brusa, M., et al. 2007, *ApJS*, **172**, 1
- Selsing, J., Fynbo, J. P. U., Christensen, L., & Krogager, J.-K. 2016, *A&A*, **585**, A87
- Sérsic, J. L. 1968, *Atlas de Galaxias Australes* (Cordoba: Observatorio Astronomico)
- Shen, S., Mo, H. J., White, S. D. M., et al. 2003, *MNRAS*, **343**, 978
- Shih, H.-Y., & Stockton, A. 2011, *ApJ*, **733**, 45
- Sparre, M. 2015, XSHPipelineManager: Wrapper for the VLT/X-shooter Data Reduction Pipeline, v 5.9.0, Astrophysics Source Code Library, ascl:1509.001
- Speagle, J. S., Steinhardt, C. L., Capak, P. L., & Silverman, J. D. 2014, *ApJS*, **214**, 15
- Szomoru, D., Franx, M., Bouwens, R. J., et al. 2011, *ApJL*, **735**, L22
- Szomoru, D., Franx, M., van Dokkum, P. G., et al. 2010, *ApJL*, **714**, L244
- Szomoru, D., Franx, M., & van Dokkum, P. G. 2012, *ApJ*, **749**, 121
- Tapia, T., Eliche-Moral, M. C., Querejeta, M., et al. 2014, *A&A*, **565**, A31
- Taylor, E. N., Franx, M., Brinchmann, J., van der Wel, A., & van Dokkum, P. G. 2010a, *ApJ*, **722**, 1
- Taylor, E. N., Franx, M., Glazebrook, K., et al. 2010b, *ApJ*, **720**, 723
- Toft, S., Franx, M., van Dokkum, P., et al. 2009, *ApJ*, **705**, 255
- Toft, S., Gallazzi, A., Zirm, A., et al. 2012, *ApJ*, **754**, 3
- Toft, S., van Dokkum, P., Franx, M., et al. 2005, *ApJL*, **624**, L9
- Toft, S., van Dokkum, P., Franx, M., et al. 2007, *ApJ*, **671**, 285
- Toft, S., Zabl, J., Richard, J., et al. 2017, *Natur*, **546**, 510
- Torrey, P., Wellons, S., Ma, C.-P., Hopkins, P. F., & Vogelsberger, M. 2017, *MNRAS*, **467**, 4872
- Trujillo, I., Cenarro, A. J., de Lorenzo-Cáceres, A., et al. 2009, *ApJL*, **692**, L118
- Trujillo, I., Conselice, C. J., Bundy, K., et al. 2007, *MNRAS*, **382**, 109
- Trujillo, I., Förster Schreiber, N. M., Rudnick, G., et al. 2006, *ApJ*, **650**, 18
- Utomo, D., Kriek, M., Labbé, I., Conroy, C., & Fumagalli, M. 2014, *ApJL*, **783**, L30
- van de Sande, J., Kriek, M., Franx, M., et al. 2013, *ApJ*, **771**, 85
- van de Sande, J., Kriek, M., Franx, M., Bezanson, R., & van Dokkum, P. G. 2014, *ApJL*, **793**, L31
- van de Sande, J., Lagos, C. D. P., Welker, C., et al. 2019, *MNRAS*, **484**, 869
- van der Wel, A., Bell, E. F., van den Bosch, F. C., Gallazzi, A., & Rix, H.-W. 2009, *ApJ*, **698**, 1232
- van der Wel, A., Franx, M., van Dokkum, P. G., et al. 2014, *ApJ*, **788**, 28
- van der Wel, A., Holden, B. P., Zirm, A. W., et al. 2008, *ApJ*, **688**, 48
- van der Wel, A., Rix, H.-W., Wuyts, S., et al. 2011, *ApJ*, **730**, 38
- van Dokkum, P. G., & Franx, M. 1996, *MNRAS*, **281**, 985
- van Dokkum, P. G., Franx, M., Kriek, M., et al. 2008, *ApJL*, **677**, L5
- van Dokkum, P. G., Quadri, R., Marchesini, D., et al. 2006, *ApJL*, **638**, L59
- van Dokkum, P. G., Whitaker, K. E., Brammer, G., et al. 2010, *ApJ*, **709**, 1018
- Veale, M., Ma, C.-P., Greene, J. E., et al. 2018, *MNRAS*, **473**, 5446
- Vernet, J., Dekker, H., D'Odorico, S., et al. 2011, *A&A*, **536**, A105
- Wellons, S., & Torrey, P. 2017, *MNRAS*, **467**, 3887
- Whitaker, K. E., Bezanson, R., van Dokkum, P. G., et al. 2017, *ApJ*, **838**, 19
- Whitaker, K. E., Labbé, I., van Dokkum, P. G., et al. 2011, *ApJ*, **735**, 86
- Williams, R. J., Quadri, R. F., Franx, M., et al. 2010, *ApJ*, **713**, 738
- Williams, R. J., Quadri, R. F., Franx, M., van Dokkum, P., & Labbé, I. 2009, *ApJ*, **691**, 1879
- Wu, H., Cao, C., Hao, C.-N., et al. 2005, *ApJL*, **632**, L79
- Wuyts, S., Förster Schreiber, N. M., van der Wel, A., et al. 2011, *ApJ*, **742**, 96
- Zabl, J., Nørgaard-Nielsen, H. U., Fynbo, J. P. U., et al. 2015, *MNRAS*, **451**, 2050
- Zahid, H. J., Geller, M. J., Damjanov, I., & Sohn, J. 2019, *ApJ*, **878**, 158
- Zamojski, M. A., Schiminovich, D., Rich, R. M., et al. 2007, *ApJS*, **172**, 468
- Zhu, Y.-N., Wu, H., Cao, C., & Li, H.-N. 2008, *ApJ*, **686**, 155
- Zibetti, S., Gallazzi, A. R., Ascasibar, Y., et al. 2017, *MNRAS*, **468**, 1902
- Zibetti, S., Gallazzi, A. R., Hirschmann, M., et al. 2019, *MNRAS*, **491**, 3